

Drift effects in cosmology

Mikołaj Korzyński (CFT PAN Warsaw)

in collaboration with

Filip Ficek (FAIS UJ Kraków) Jarosław Kopiński (FUW Warsaw)

CosmoToruń 17 Inhomogeneous Cosmology Workshop

Centre for Astronomy at Nicolaus Copernicus University



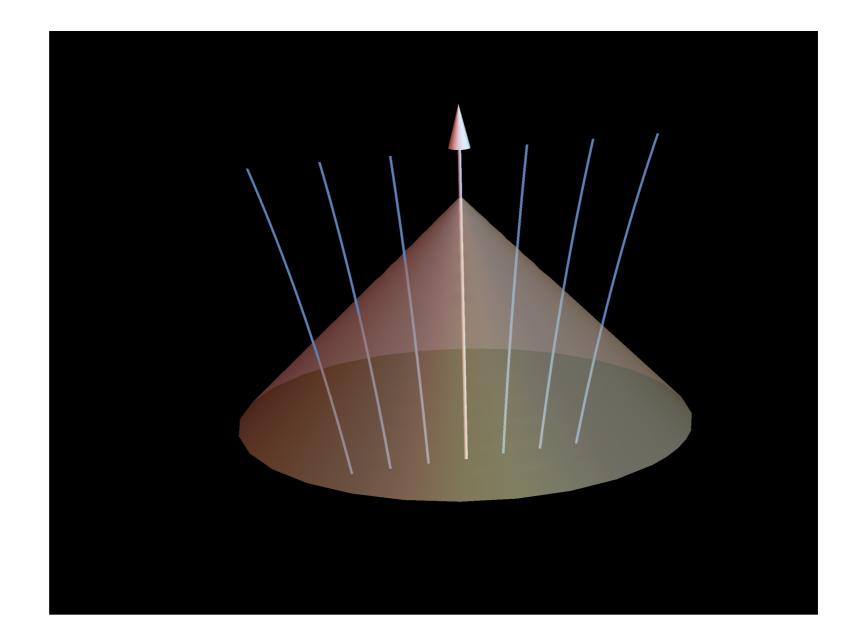
Toruń, 2nd-7th July 2017

Outline

- I. Motivation
- 2. Geometry
- 3. Drift effects
- 4. Applications

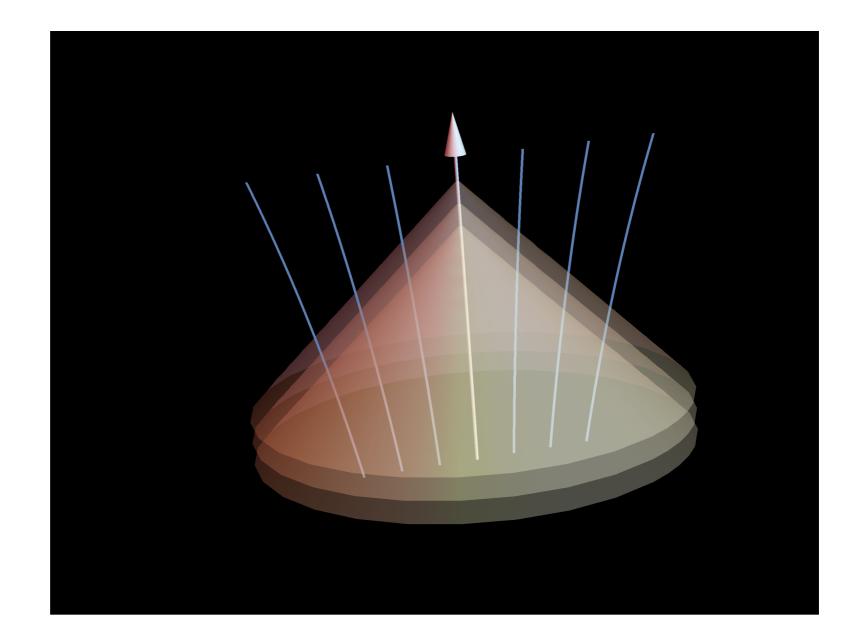
NCN project SONATA BIS No 2016/22/E/ST9/00578 "Local relativistic perturbative framework in hydrodynamics and general relativity and its application to cosmology"

Observations in cosmology effectively on a single lightcone



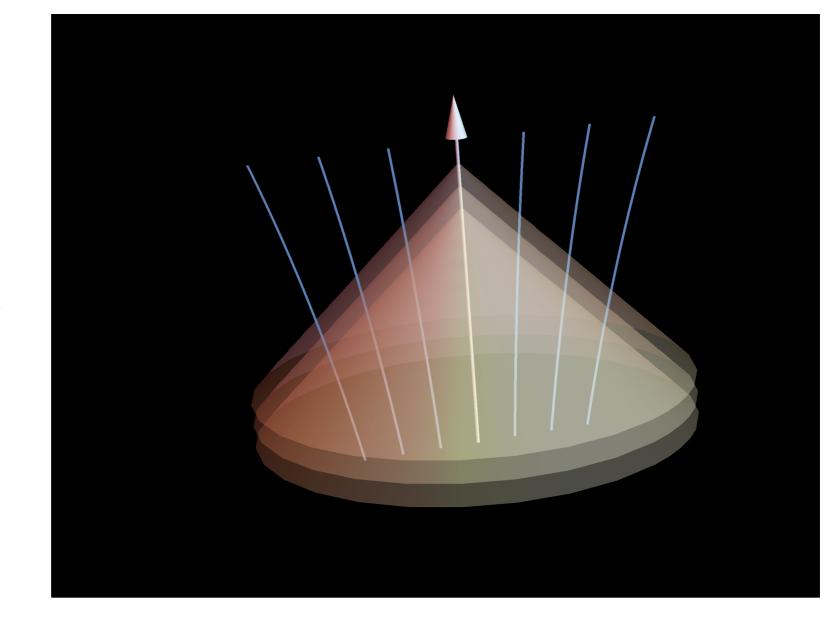
CosmoTorun I7

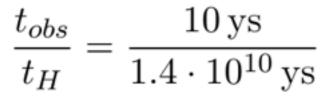
Observations in cosmology effectively on a single lightcone



CosmoTorun I7

Observations in cosmology effectively on a single lightcone





CosmoTorun I7

In FLRW things are quite simple (non-rotating frame, cosmic flow observer and emitter)

In FLRW things are quite simple (non-rotating frame, cosmic flow observer and emitter)

$$\frac{\mathrm{d}}{\mathrm{d}t}\theta^{A} = 0$$

$$\frac{\mathrm{d}}{\mathrm{d}t}\ln(1+z) = H(t_{obs}) - \frac{1}{1+z}H(t_{em})$$

$$\frac{\mathrm{d}}{\mathrm{d}t}\ln D_{area} = H(t_{obs}) - \frac{1}{1+z}H(t_{em})$$

$$\frac{\mathrm{d}}{\mathrm{d}t}\ln D_{lum} = 3\left(H(t_{obs}) - \frac{1}{1+z}H(t_{em})\right)$$

$$\frac{D_{area}}{1+z} = \mathrm{const}$$

In FLRW things are quite simple (non-rotating frame, cosmic flow observer and emitter)

$$\frac{d}{dt}\theta^{A} = 0$$

$$\frac{d}{dt}\ln(1+z) = H(t_{obs}) - \frac{1}{1+z}H(t_{em})$$
Sandage 1962
$$\frac{d}{dt}\ln D_{area} = H(t_{obs}) - \frac{1}{1+z}H(t_{em})$$

$$\frac{d}{dt}\ln D_{lum} = 3\left(H(t_{obs}) - \frac{1}{1+z}H(t_{em})\right)$$

$$\frac{D_{area}}{1+z} = \text{const}$$

In FLRW things are quite simple (non-rotating frame, cosmic flow observer and emitter)

$$\frac{d}{dt}\theta^{A} = 0$$

$$\frac{d}{dt}\ln(1+z) = H(t_{obs}) - \frac{1}{1+z}H(t_{em})$$
Sandage 1962
$$\frac{d}{dt}\ln D_{area} = H(t_{obs}) - \frac{1}{1+z}H(t_{em})$$

$$\frac{d}{dt}\ln D_{lum} = 3\left(H(t_{obs}) - \frac{1}{1+z}H(t_{em})\right)$$

$$\frac{D_{area}}{1+z} = \text{const}$$

Possible to measure directly H(z)

In FLRW things are quite simple (non-rotating frame, cosmic flow observer and emitter)

$$\frac{d}{dt}\theta^{A} = 0$$

$$\frac{d}{dt}\ln(1+z) = H(t_{obs}) - \frac{1}{1+z}H(t_{em})$$
Sandage 1962
$$\frac{d}{dt}\ln D_{area} = H(t_{obs}) - \frac{1}{1+z}H(t_{em})$$

$$\frac{d}{dt}\ln D_{lum} = 3\left(H(t_{obs}) - \frac{1}{1+z}H(t_{em})\right)$$

$$\frac{D_{area}}{1+z} = \text{const}$$

Possible to measure directly H(z)

In practice: peculiar motions, inhomogeneities (lensing, timedependent light bending), time-dependent potential wells ...

Feasibility of observations

Feasibility of observations

- E-ELT, CODEX spectrograph (Ly- α forest z drift, after 1-2 decades of observations)
- VLT ESPRESSO spectrograph
- SKA z drift in neutral hydrogen
- Gaia high accuracy position drift measurements (quasars)

Feasibility of observations

- E-ELT, CODEX spectrograph (Ly- α forest z drift, after 1-2 decades of observations)
- VLT ESPRESSO spectrograph
- SKA z drift in neutral hydrogen
- Gaia high accuracy position drift measurements (quasars)

Quercellini et al, "Real time cosmology" Physics Reports 521 (2012) 95-134

Feasibility of observations

- E-ELT, CODEX spectrograph (Ly- α forest z drift, after 1-2 decades of observations)
- VLT ESPRESSO spectrograph
- SKA z drift in neutral hydrogen
- Gaia high accuracy position drift measurements (quasars)

Quercellini et al, "Real time cosmology" Physics Reports 521 (2012) 95-134

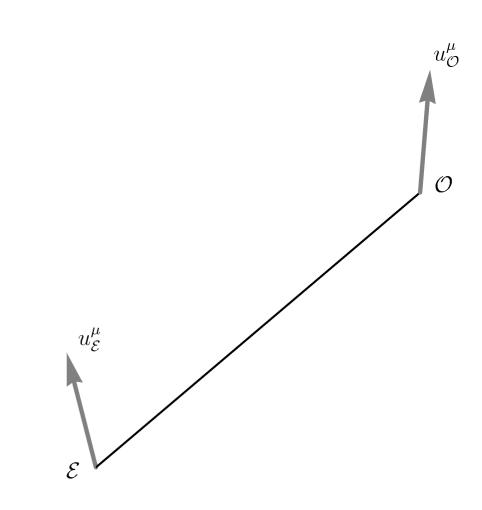
Theory

very simplified: Loeb 1998, Balbi, Quercellini 2007, Uzan et al 2008...

exact models: Krasiński 2011, Krasiński, Bolejko 2012, Quercellini et al 2009...

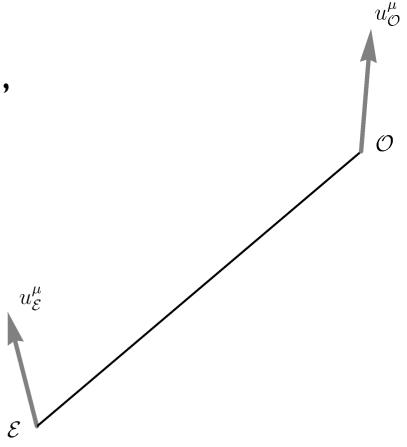
CosmoTorun I 7

Lack of a general theory



Lack of a general theory

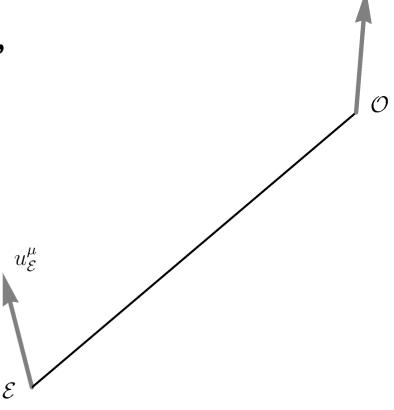
Sachs formalism (Sachs 1961, Etherington, Trautman, ...), V. Perlick Living Reviews in Relativity 7 (9) 2004



Lack of a general theory

Sachs formalism (Sachs 1961, Etherington, Trautman, ...), V. Perlick Living Reviews in Relativity 7 (9) 2004

• Provides *z*, D_{ang} , D_{lum} and $\mathcal{D}^{4}{}_{B}$ given a null geodesic connecting \mathcal{E} and O and geometry in a narrow tube around it ($R^{\mu}{}_{\nu\alpha\beta}$)

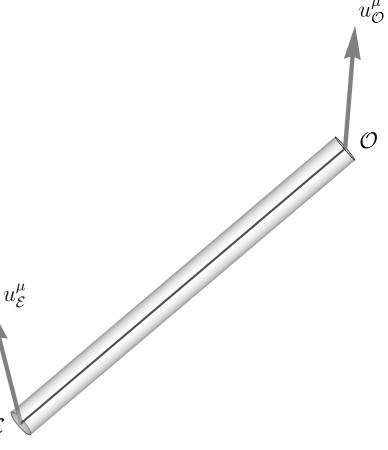


 $u^{\mu}_{\mathcal{O}}$

Lack of a general theory

Sachs formalism (Sachs 1961, Etherington, Trautman, ...), V. Perlick Living Reviews in Relativity 7 (9) 2004

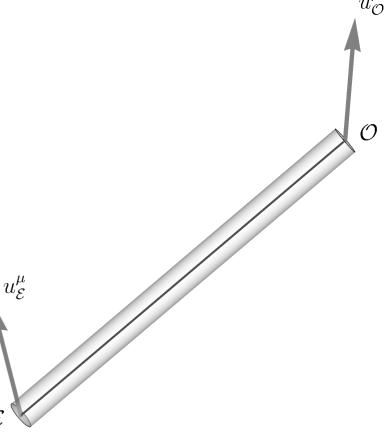
• Provides *z*, D_{ang} , D_{lum} and $\mathcal{D}^{A}{}_{B}$ given a null geodesic connecting \mathcal{E} and O and geometry in a narrow tube around it ($R^{\mu}{}_{\nu\alpha\beta}$)



Lack of a general theory

Sachs formalism (Sachs 1961, Etherington, Trautman, ...), V. Perlick Living Reviews in Relativity 7 (9) 2004

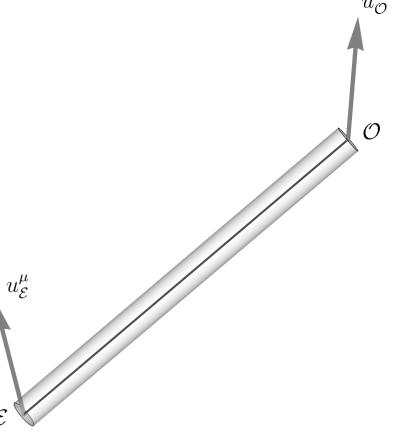
- Provides *z*, D_{ang} , D_{lum} and $\mathcal{D}^{4}{}_{B}$ given a null geodesic connecting \mathcal{E} and O and geometry in a narrow tube around it ($R^{\mu}{}_{\nu\alpha\beta}$)
- Based on the geodesic deviation equation (GDE)



Lack of a general theory

Sachs formalism (Sachs 1961, Etherington, Trautman, ...), V. Perlick Living Reviews in Relativity 7 (9) 2004

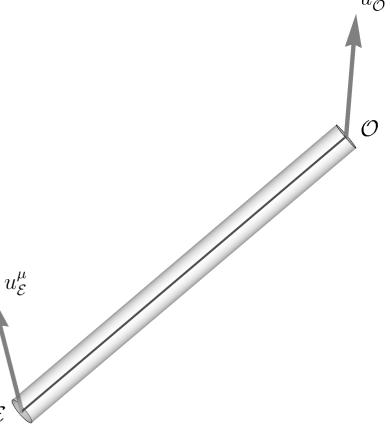
- Provides *z*, D_{ang} , D_{lum} and $\mathcal{D}^{A}{}_{B}$ given a null geodesic connecting \mathcal{E} and O and geometry in a narrow tube around it ($R^{\mu}{}_{\nu\alpha\beta}$)
- Based on the geodesic deviation equation (GDE)
- Exact (all GR effects, all spacetimes)



Lack of a general theory

Sachs formalism (Sachs 1961, Etherington, Trautman, ...), V. Perlick Living Reviews in Relativity 7 (9) 2004

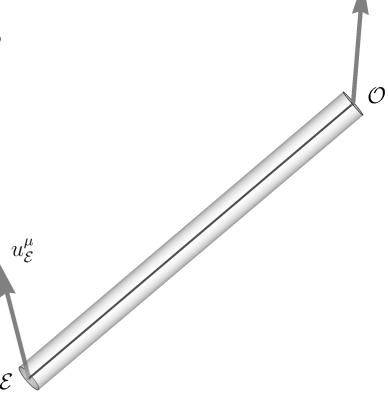
- Provides *z*, D_{ang} , D_{lum} and $\mathcal{D}^{A}{}_{B}$ given a null geodesic connecting \mathcal{E} and O and geometry in a narrow tube around it ($R^{\mu}{}_{\nu\alpha\beta}$)
- Based on the geodesic deviation equation (GDE)
- Exact (all GR effects, all spacetimes)
- Limitations: geometric optics, narrow beam



Lack of a general theory

Sachs formalism (Sachs 1961, Etherington, Trautman, ...), V. Perlick Living Reviews in Relativity 7 (9) 2004

- Provides *z*, D_{ang} , D_{lum} and $\mathcal{D}^{A}{}_{B}$ given a null geodesic connecting \mathcal{E} and O and geometry in a narrow tube around it ($R^{\mu}{}_{\nu\alpha\beta}$)
- Based on the geodesic deviation equation (GDE)
- Exact (all GR effects, all spacetimes)
- Limitations: geometric optics, narrow beam
- Applications: numerics, gravitational lensing, (weak lensing...)



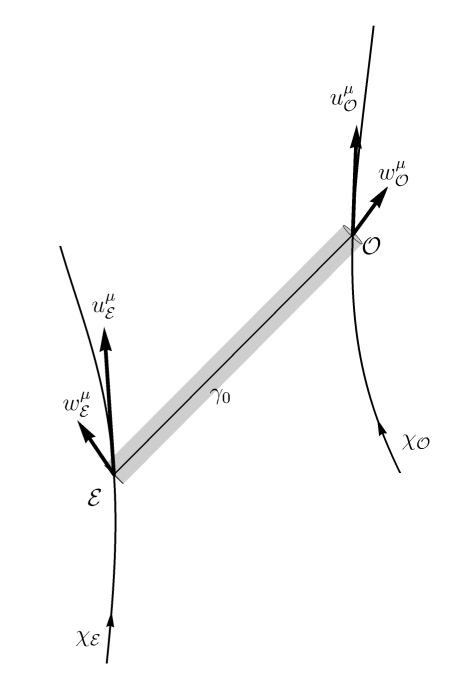
CosmoTorun I7

M. Korzyński, "Drift effects in cosmology"

Drift effects in GR - extending Sachs formalism

Drift effects in GR - extending Sachs formalism

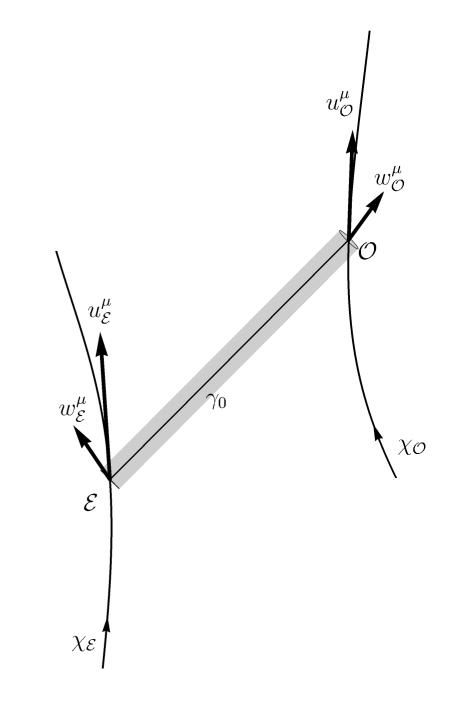
$$\begin{aligned} \frac{\mathrm{d}}{\mathrm{d}\tau} \theta^{A} &\equiv \dot{\theta}^{A} \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{\mathrm{d}}{\mathrm{d}\tau} \ln(1+z) &\equiv (1+z)^{-1} \dot{z} \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{\mathrm{d}}{\mathrm{d}\tau} \mathcal{D}^{A}{}_{B} &\equiv \dot{\mathcal{D}}^{A}{}_{B} \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{\mathrm{d}}{\mathrm{d}\tau} \ln D_{area} &\equiv (\ln D_{area}) \cdot \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{\mathrm{d}}{\mathrm{d}\tau} \ln D_{lum} &\equiv (\ln D_{lum}) \cdot \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \end{aligned}$$



Drift effects in GR - extending Sachs formalism

$$\begin{split} &\frac{\mathrm{d}}{\mathrm{d}\tau}\theta^{A} \equiv \dot{\theta}^{A} \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}_{\ \nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \ln(1+z) \equiv (1+z)^{-1} \dot{z} \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}_{\ \nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \mathcal{D}^{A}{}_{B} \equiv \dot{\mathcal{D}}^{A}{}_{B} \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \ln D_{area} \equiv (\ln D_{area}) \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \ln D_{lum} \equiv (\ln D_{lum}) \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta}\right) \end{split}$$

Based on the geodesic deviation equation (GDE)

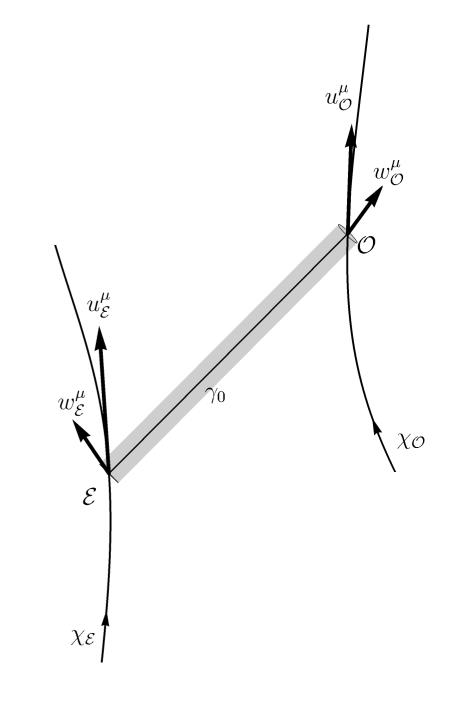


CosmoTorun I7

Drift effects in GR - extending Sachs formalism

$$\begin{split} &\frac{\mathrm{d}}{\mathrm{d}\tau}\theta^{A} \equiv \dot{\theta}^{A} \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}_{\ \nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \ln(1+z) \equiv (1+z)^{-1} \dot{z} \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}_{\ \nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \mathcal{D}^{A}{}_{B} \equiv \dot{\mathcal{D}}^{A}{}_{B} \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \ln D_{area} \equiv (\ln D_{area}) \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \ln D_{lum} \equiv (\ln D_{lum}) \left(\mathcal{O}, \mathcal{E}, u^{\mu}_{\mathcal{O}}, u^{\mu}_{\mathcal{E}}, w^{\mu}_{\mathcal{O}}, w^{\mu}_{\mathcal{E}}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta}\right) \end{split}$$

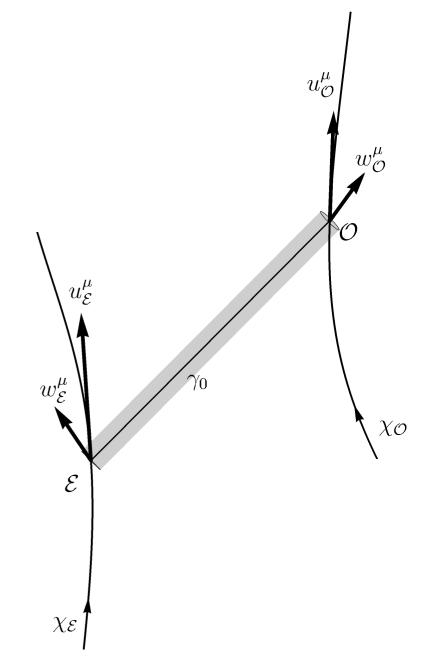
- Based on the geodesic deviation equation (GDE)
- Exact (all GR effects, all spacetimes)



Drift effects in GR - extending Sachs formalism

$$\begin{split} &\frac{\mathrm{d}}{\mathrm{d}\tau}\theta^{A} \equiv \dot{\theta}^{A} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}_{\ \nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \ln(1+z) \equiv (1+z)^{-1} \dot{z} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}_{\ \nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \mathcal{D}^{A}{}_{B} \equiv \dot{\mathcal{D}}^{A}{}_{B} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \ln D_{area} \equiv (\ln D_{area}) \dot{} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta}\right) \\ &\frac{\mathrm{d}}{\mathrm{d}\tau} \ln D_{lum} \equiv (\ln D_{lum}) \dot{} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta}\right) \end{split}$$

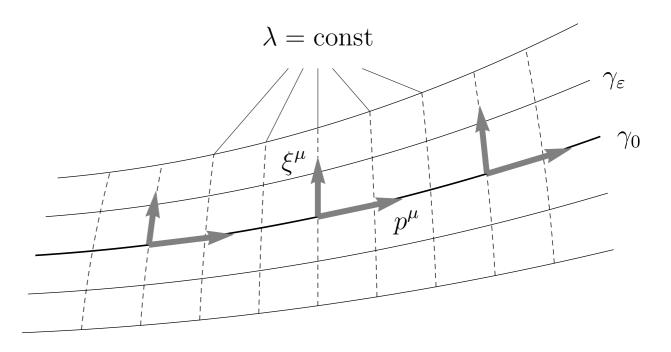
- Based on the geodesic deviation equation (GDE)
- Exact (all GR effects, all spacetimes)
- Local effects (observer, emitter) vs propagation effects



CosmoTorun I7

Geodesic deviation equation

$$\mathcal{G}[\xi]^{\mu} \equiv \nabla_p \nabla_p \xi^{\mu} - R^{\mu}_{\ \nu\alpha\beta} \, p^{\nu} \, p^{\alpha} \, \xi^{\beta} = 0$$

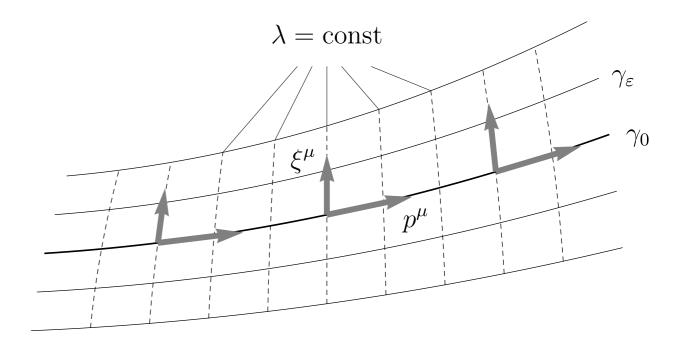


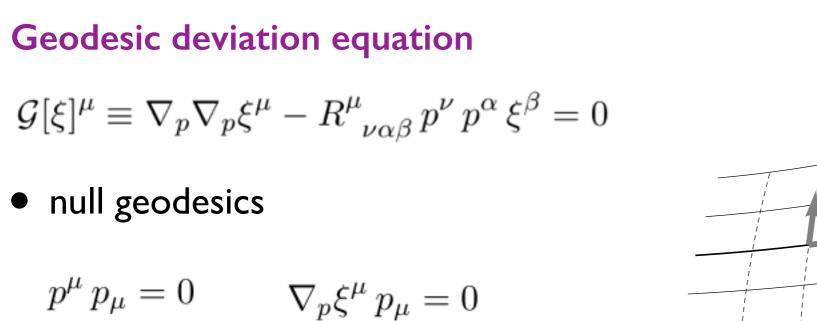
Geodesic deviation equation

$$\mathcal{G}[\xi]^{\mu} \equiv \nabla_p \nabla_p \xi^{\mu} - R^{\mu}_{\ \nu\alpha\beta} \, p^{\nu} \, p^{\alpha} \, \xi^{\beta} = 0$$

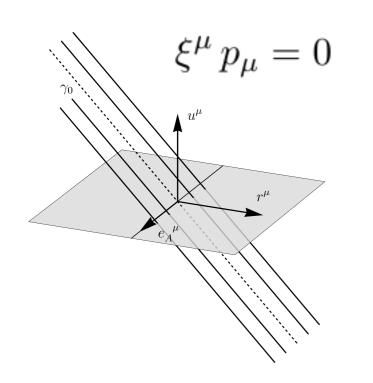
• null geodesics

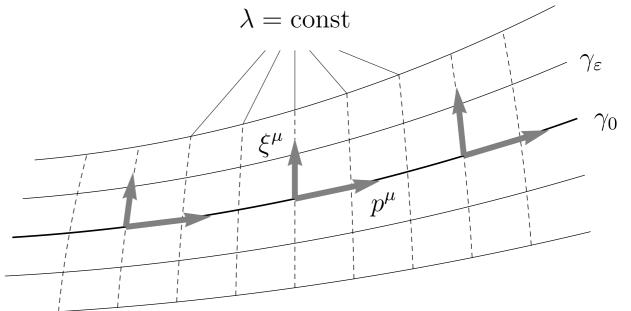
$$p^{\mu} p_{\mu} = 0 \qquad \nabla_p \xi^{\mu} p_{\mu} = 0$$

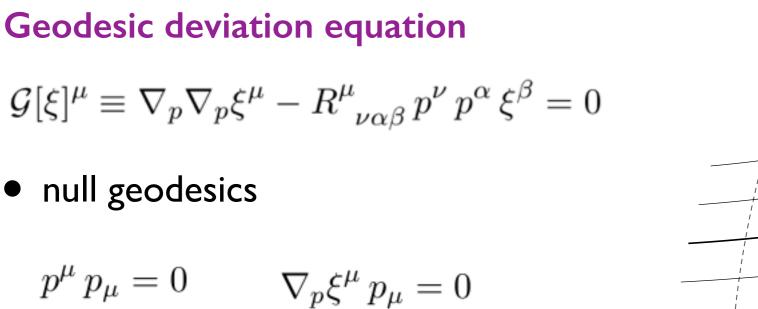




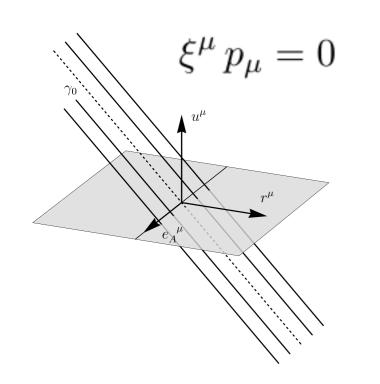
• special case: orthogonally displaced null geodesics



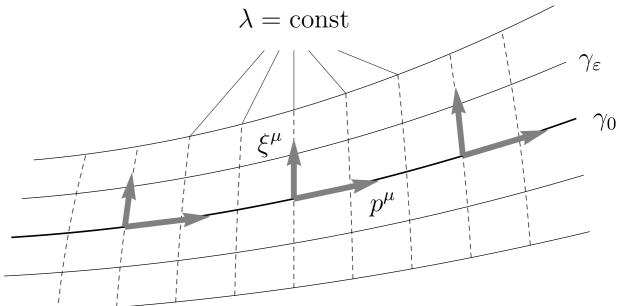


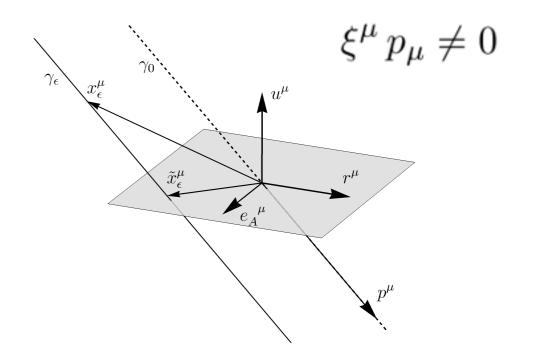






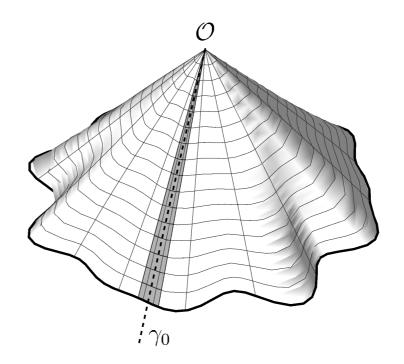
CosmoTorun I7





Jacobi matrix

$$\xi^{A}(\lambda) = \mathcal{D}^{A}{}_{B}(\lambda) \,\nabla_{p} \xi^{A}(\lambda_{\mathcal{O}})$$

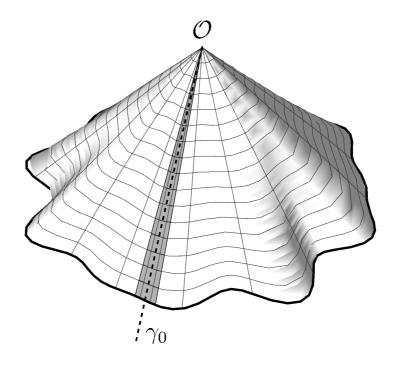


Jacobi matrix

$$\xi^{A}(\lambda) = \mathcal{D}^{A}{}_{B}(\lambda) \,\nabla_{p} \xi^{A}(\lambda_{\mathcal{O}})$$

• given by ODE's

$$\frac{\mathrm{d}^2}{\mathrm{d}\lambda^2} \mathcal{D}^A{}_B - R^A{}_{\nu\alpha C} p^{\nu} p^{\alpha} \mathcal{D}^C{}_B = 0$$
$$\mathcal{D}^A{}_B (\lambda_{\mathcal{O}}) = 0$$
$$\frac{\mathrm{d}}{\mathrm{d}\lambda} \mathcal{D}^A{}_B (\lambda_{\mathcal{O}}) = \delta^A{}_B$$



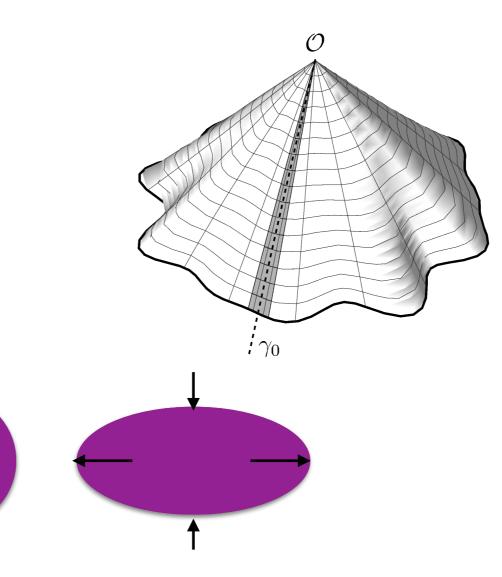
Jacobi matrix

$$\xi^{A}(\lambda) = \mathcal{D}^{A}{}_{B}(\lambda) \,\nabla_{p} \xi^{A}(\lambda_{\mathcal{O}})$$

• given by ODE's

$$\frac{\mathrm{d}^2}{\mathrm{d}\lambda^2} \mathcal{D}^A{}_B - R^A{}_{\nu\alpha C} p^{\nu} p^{\alpha} \mathcal{D}^C{}_B = 0$$
$$\mathcal{D}^A{}_B (\lambda_{\mathcal{O}}) = 0$$
$$\frac{\mathrm{d}}{\mathrm{d}\lambda} \mathcal{D}^A{}_B (\lambda_{\mathcal{O}}) = \delta^A{}_B$$

• gravitational lensing

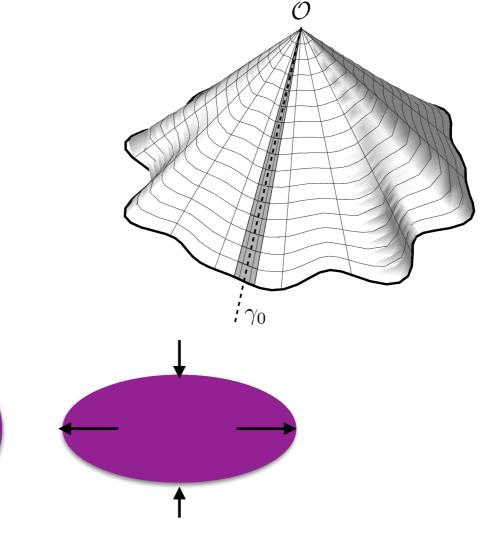


Jacobi matrix

$$\xi^{A}(\lambda) = \mathcal{D}^{A}{}_{B}(\lambda) \, \nabla_{p} \xi^{A}(\lambda_{\mathcal{O}})$$

• given by ODE's

$$\frac{\mathrm{d}^2}{\mathrm{d}\lambda^2} \mathcal{D}^A{}_B - R^A{}_{\nu\alpha C} p^{\nu} p^{\alpha} \mathcal{D}^C{}_B = 0$$
$$\mathcal{D}^A{}_B (\lambda_{\mathcal{O}}) = 0$$
$$\frac{\mathrm{d}}{\mathrm{d}\lambda} \mathcal{D}^A{}_B (\lambda_{\mathcal{O}}) = \delta^A{}_B$$



- gravitational lensing
- area, luminosity distances

$$D_{ang} = \left(p_{\mu} \, u_{\mathcal{O}}^{\mu} \right) \, \left| \det \mathcal{D}^{A}{}_{B} \left(\lambda_{\mathcal{E}} \right) \right|^{1/2}$$

 $D_{lum} = D_{ang}(1+z)^2$

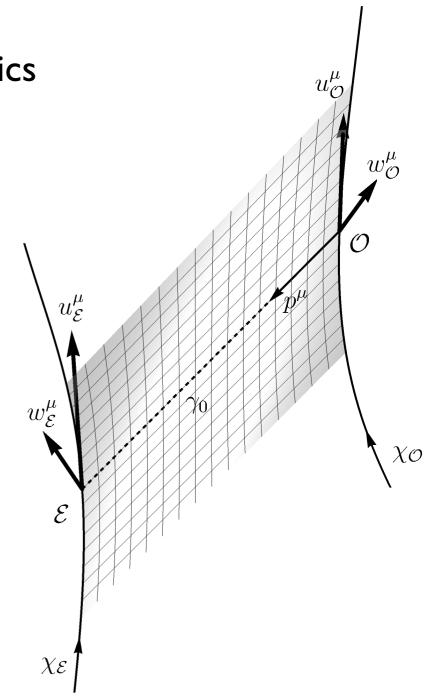
CosmoTorun I7



Geometric setup

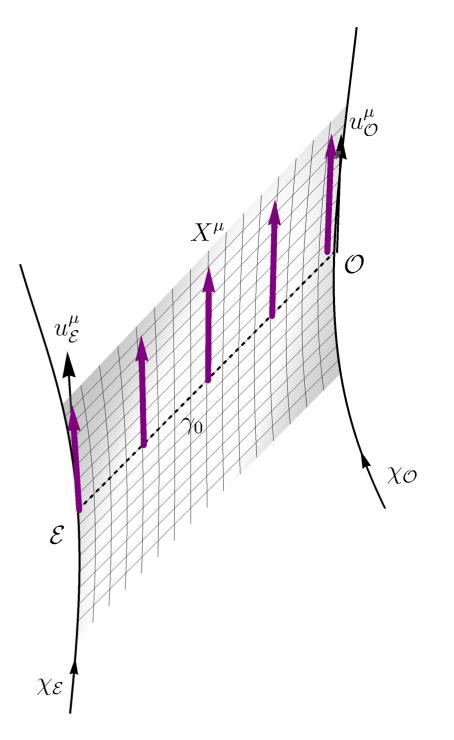
Geometry

• null surface spanned by connecting null geodesics



Geometry

- null surface spanned by connecting null geodes
- main tool: observation time vector X^{μ}



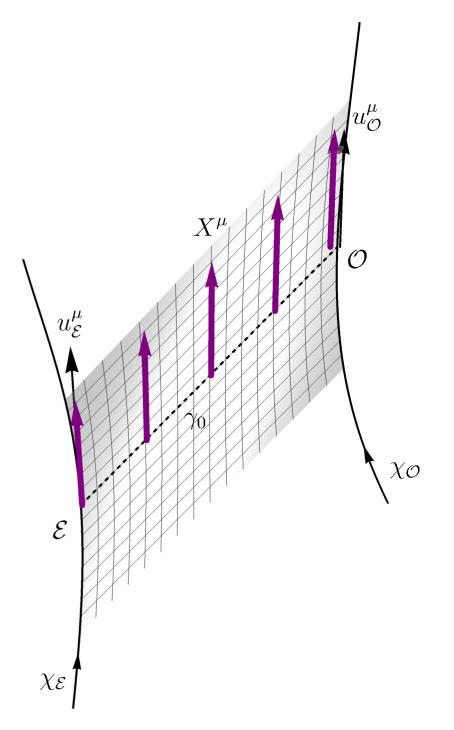
Geometry

- null surface spanned by connecting null geodes
- main tool: observation time vector X^{μ}

$$\mathcal{G}[X]^{\mu} = 0$$

$$X^{\mu}(\lambda_{\mathcal{O}}) = u^{\mu}_{\mathcal{O}}$$

$$X^{\mu}(\lambda_{\mathcal{E}}) = \frac{1}{1+z} u^{\mu}_{\mathcal{E}}$$



Geometry

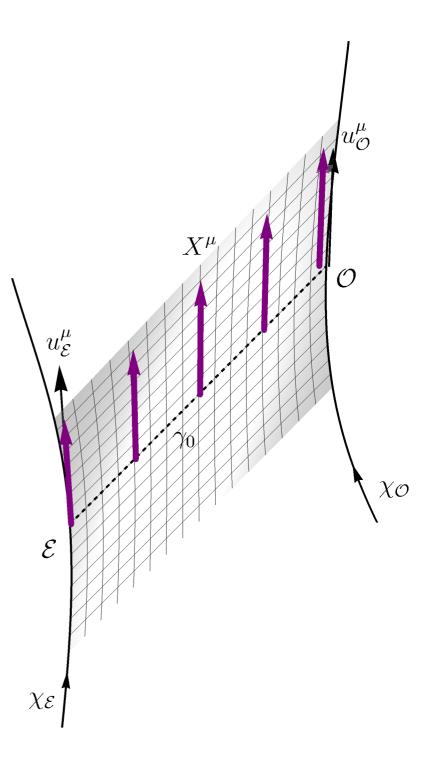
- null surface spanned by connecting null geodes
- main tool: observation time vector X^{μ}

$$\mathcal{G}[X]^{\mu} = 0$$

$$X^{\mu}(\lambda_{\mathcal{O}}) = u^{\mu}_{\mathcal{O}}$$

$$X^{\mu}(\lambda_{\mathcal{E}}) = \frac{1}{1+z} u^{\mu}_{\mathcal{E}}$$

• can be solved (up to $C \cdot p^{\mu}$) using $\mathcal{D}^{A_{B}}$



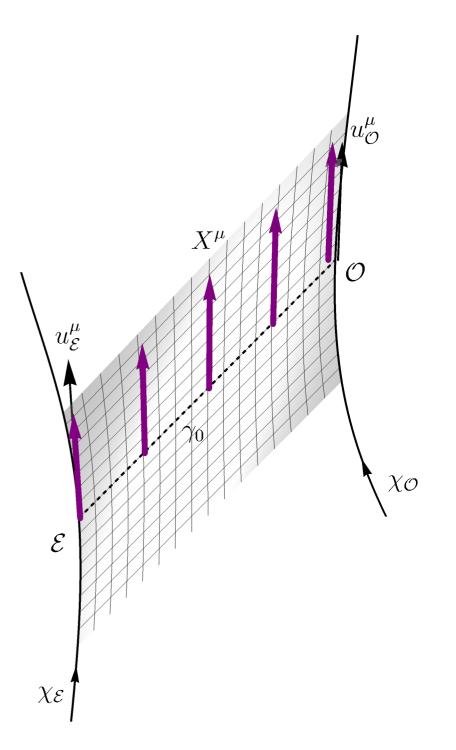
Geometry

- null surface spanned by connecting null geodes
- main tool: observation time vector X^{μ}

$$\begin{aligned} X^{\mu}(\lambda_{\mathcal{O}}) &= u^{\mu}_{\mathcal{O}} \\ X^{\mu}(\lambda_{\mathcal{E}}) &= \frac{1}{1+z} u^{\mu}_{\mathcal{E}} \end{aligned}$$

 $\mathcal{G}[X]^{\mu} = 0$

- can be solved (up to $C \cdot p^{\mu}$) using $\mathcal{D}^{A_{B}}$
- ∇_X gives the drift effects



Geometry

- null surface spanned by connecting null geodes
- main tool: observation time vector X^{μ}

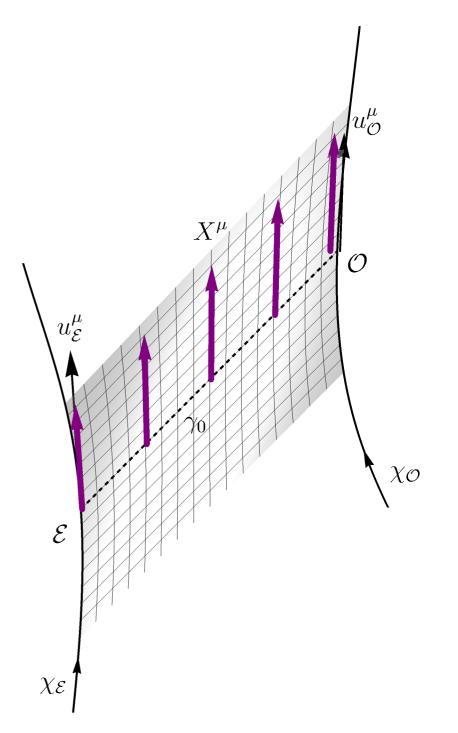
$$\mathcal{G}[X]^{\mu} = 0$$

$$X^{\mu}(\lambda_{\mathcal{O}}) = u^{\mu}_{\mathcal{O}}$$

$$X^{\mu}(\lambda_{\mathcal{E}}) = \frac{1}{1+z} u^{\mu}_{\mathcal{E}}$$

- can be solved (up to $C \cdot p^{\mu}$) using $\mathcal{D}^{A_{B}}$
- ∇_X gives the drift effects

 $abla_X p^{\mu}$ - position drift



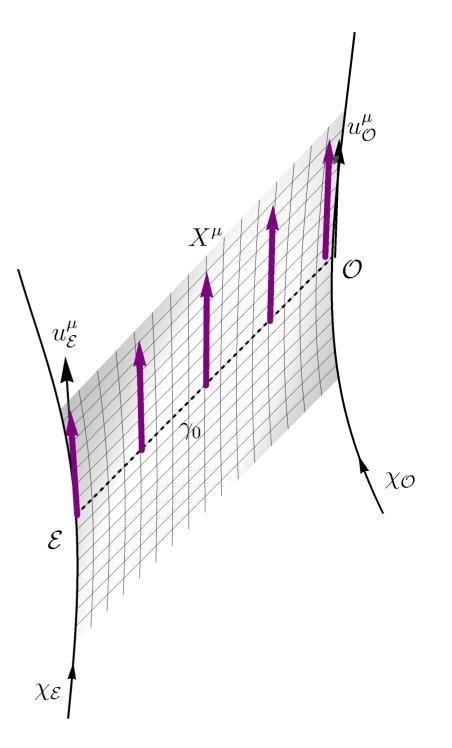
Geometry

- null surface spanned by connecting null geodes
- main tool: observation time vector X^{μ}

$$\begin{aligned} & \mathcal{J}^{\mu}(\lambda_{\mathcal{O}}) &= u_{\mathcal{O}}^{\mu} \\ & X^{\mu}(\lambda_{\mathcal{E}}) &= \frac{1}{1+z} u_{\mathcal{E}}^{\mu} \end{aligned}$$

 $\mathcal{G}[X]^{\mu} = 0$

- can be solved (up to $C \cdot p^{\mu}$) using $\mathcal{D}^{A_{B}}$
- ∇_X gives the drift effects $\nabla_X p^{\mu}$ - position drift $\nabla_X (p_{\mu} u^{\mu})$ - redshift drift



Geometry

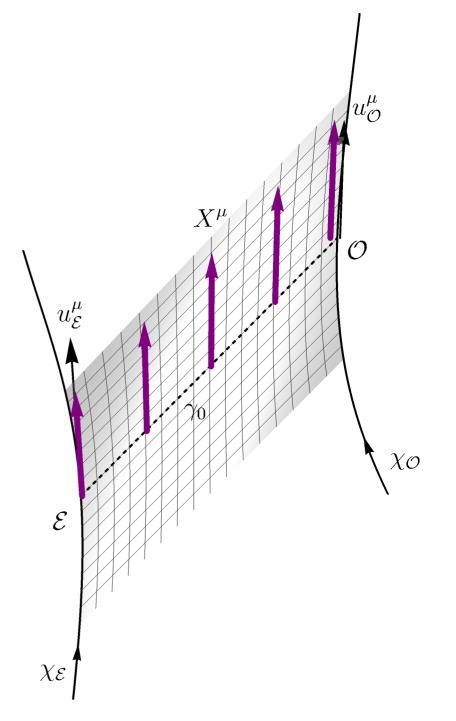
- null surface spanned by connecting null geodes
- main tool: observation time vector X^{μ}

$$\mathcal{G}[X]^{\mu} = 0$$

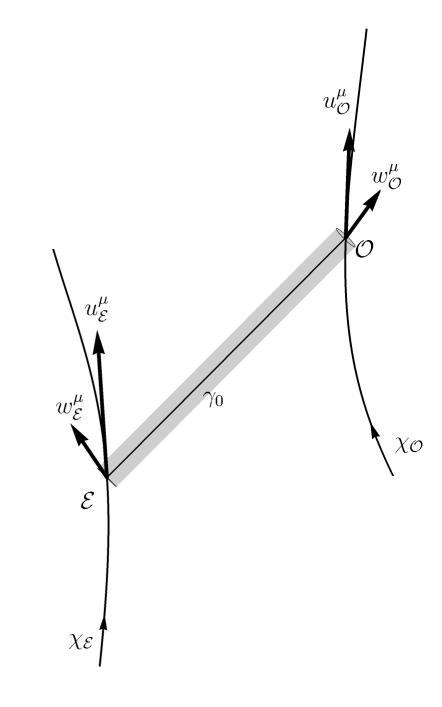
$$X^{\mu}(\lambda_{\mathcal{O}}) = u^{\mu}_{\mathcal{O}}$$

$$X^{\mu}(\lambda_{\mathcal{E}}) = \frac{1}{1+z} u^{\mu}_{\mathcal{E}}$$

- can be solved (up to $C \cdot p^{\mu}$) using $\mathcal{D}^{A_{B}}$
- ∇_X gives the drift effects $\nabla_X p^{\mu}$ - position drift $\nabla_X (p_{\mu} u^{\mu})$ - redshift drift $\nabla_X \mathcal{D}^A_B$ - distances drift



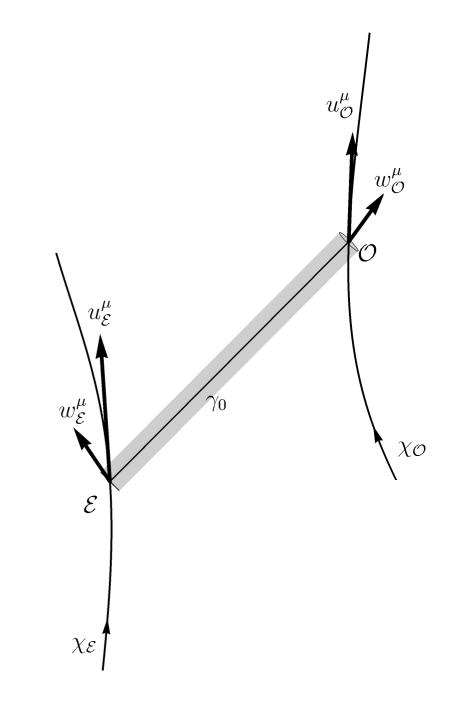
Position drift



CosmoTorun I7

Position drift

• parallel propagation of $\hat{u}^{\mu}_{\mathcal{O}}$

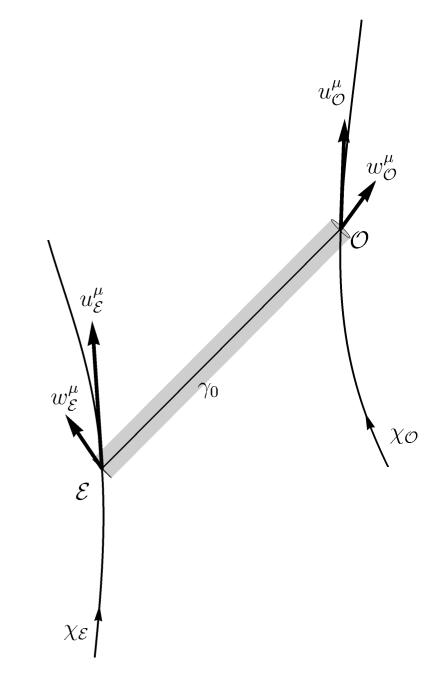


CosmoTorun I7

Position drift

- parallel propagation of $\hat{u}^{\mu}_{\mathcal{O}}$
- inhomogeneous (perpendicular) GDE

 $\begin{aligned} \mathcal{G}[m]^A &= R^A_{\ \nu\alpha\beta} \, p^\nu \, p^\alpha \, \hat{u}^\beta_{\mathcal{O}} \\ m^A(\lambda_{\mathcal{O}}) &= 0 \\ \nabla_p m^A(\lambda_{\mathcal{O}}) &= 0 \end{aligned}$



Position drift

- parallel propagation of $\hat{u}^{\mu}_{\mathcal{O}}$
- inhomogeneous (perpendicular) GDE

$$\begin{aligned} \mathcal{G}[m]^A &= R^A_{\ \nu\alpha\beta} \, p^\nu \, p^\alpha \, \hat{u}^\beta_{\mathcal{O}} \\ m^A(\lambda_{\mathcal{O}}) &= 0 \\ \nabla_p m^A(\lambda_{\mathcal{O}}) &= 0 \end{aligned}$$

$$\frac{\mathcal{D}^{F-W}}{\mathrm{d}\tau}r^A = w_{\mathcal{O}}^A + \frac{1}{p_{\sigma} u_{\mathcal{O}}^{\sigma}} \mathcal{D}^{-1} (\lambda_{\mathcal{E}})^A{}_B \left(\left(\frac{1}{1+z} u_{\mathcal{E}} - \hat{u}_{\mathcal{O}} \right)^B - m^B \right) \chi_{\mathcal{E}}$$

CosmoTorun I7

 $u_{\mathcal{E}}^{\mu}$

 $w^{\mu}_{\mathcal{E}}$

 \mathcal{E}

 $u^{\mu}_{\mathcal{O}}$

 $w^{\mu}_{\mathcal{C}}$

 $\chi_{\mathcal{O}}$

Position drift

- parallel propagation of $\hat{u}^{\mu}_{\mathcal{O}}$
- inhomogeneous (perpendicular) GDE

$$\begin{aligned} \mathcal{G}[m]^A &= R^A_{\ \nu\alpha\beta} \, p^\nu \, p^\alpha \, \hat{u}^\beta_{\mathcal{O}} \\ m^A(\lambda_{\mathcal{O}}) &= 0 \\ \nabla_p m^A(\lambda_{\mathcal{O}}) &= 0 \end{aligned}$$

٦

$$\frac{\mathbf{D}^{F-W}}{\mathrm{d}\tau}\mathbf{r}^{A} = w_{\mathcal{O}}^{A} + \frac{1}{p_{\sigma} u_{\mathcal{O}}^{\sigma}} \mathcal{D}^{-1} (\lambda_{\mathcal{E}})^{A}{}_{B} \left(\left(\frac{1}{1+z} u_{\mathcal{E}} - \hat{u}_{\mathcal{O}} \right)^{B} - m^{B} \right)_{\chi_{\mathcal{E}}}$$

mi-Walker
erivative
bserver)

CosmoTorun I 7

Ferr

d

 $\mathbf{0}$

M. Korzyński, "Drift effects in cosmology"

 $u_{\mathcal{E}}^{\mu}$

 $w^{\mu}_{\mathcal{E}}$

 \mathcal{E}

 $u^{\mu}_{\mathcal{O}}$

 $w^{\mu}_{\mathcal{O}}$

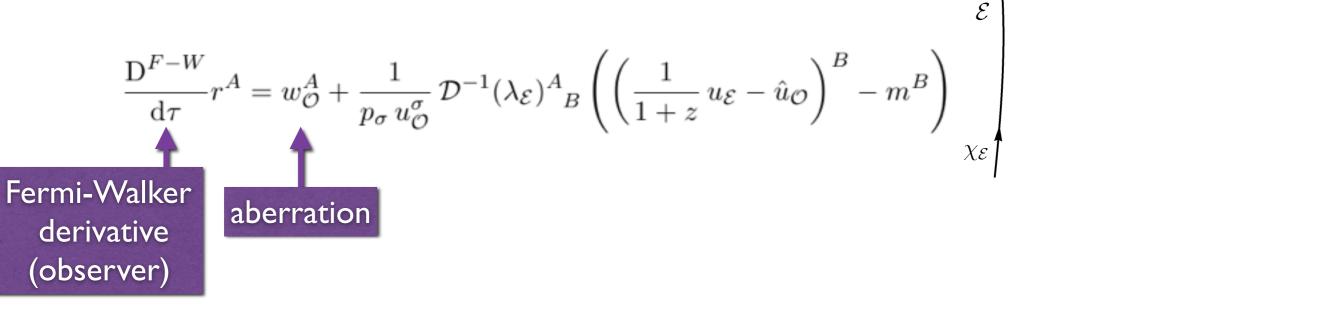
 $\chi_{\mathcal{O}}$

Position drift

- parallel propagation of $\hat{u}^{\mu}_{\mathcal{O}}$
- inhomogeneous (perpendicular) GDE

$$\begin{aligned} \mathcal{G}[m]^A &= R^A_{\ \nu\alpha\beta} \, p^\nu \, p^\alpha \, \hat{u}^\beta_{\mathcal{O}} \\ m^A(\lambda_{\mathcal{O}}) &= 0 \\ \nabla_p m^A(\lambda_{\mathcal{O}}) &= 0 \end{aligned}$$

• result



CosmoTorun 17

M. Korzyński, "Drift effects in cosmology"

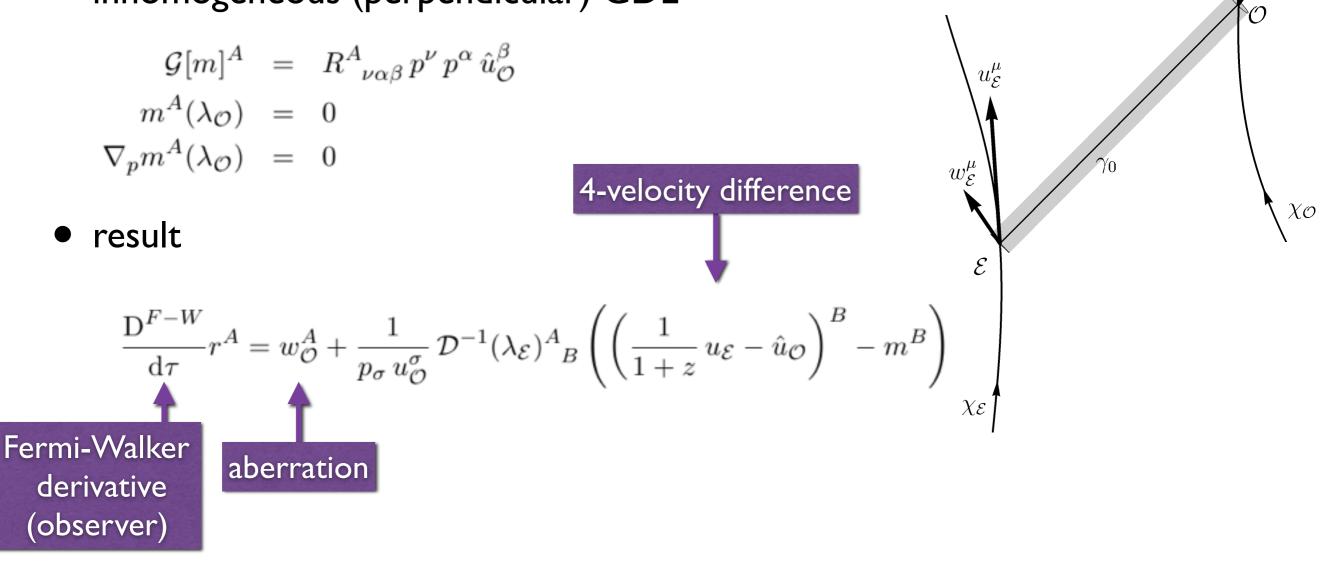
 $u_{\mathcal{E}}^{\mu}$

 $w_{\mathcal{E}}^{\mu}$

 $u^{\mu}_{\mathcal{O}}$

Position drift

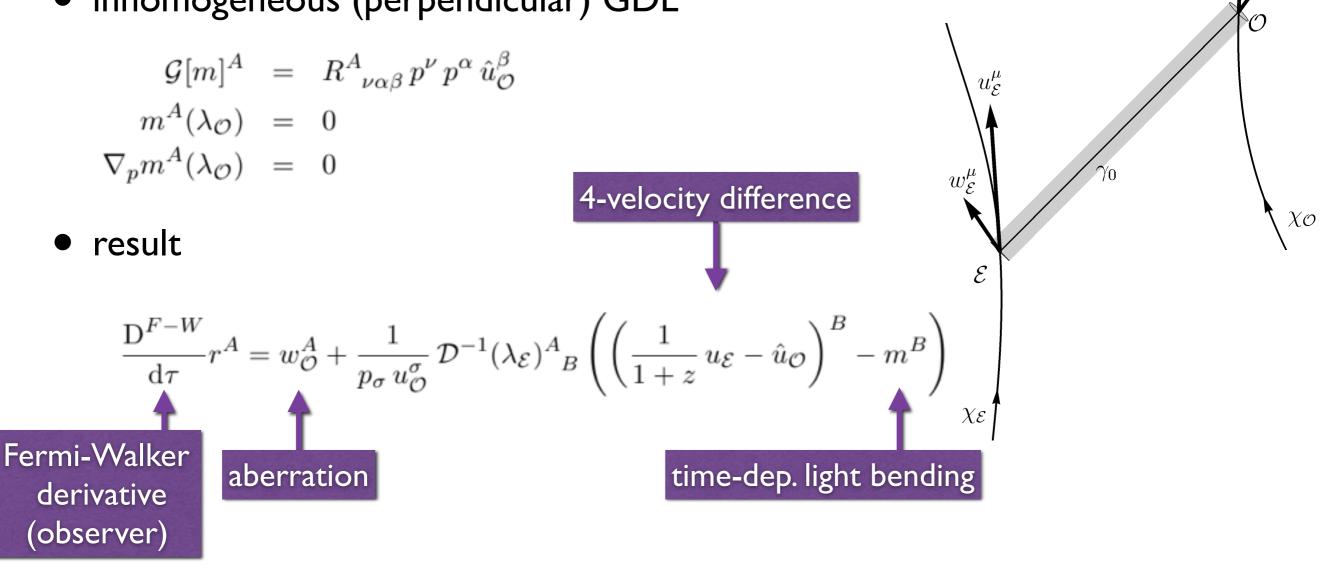
- parallel propagation of $\hat{u}^{\mu}_{\mathcal{O}}$
- inhomogeneous (perpendicular) GDE



CosmoTorun 17

Position drift

- parallel propagation of \hat{u}^{μ}_{O}
- inhomogeneous (perpendicular) GDE



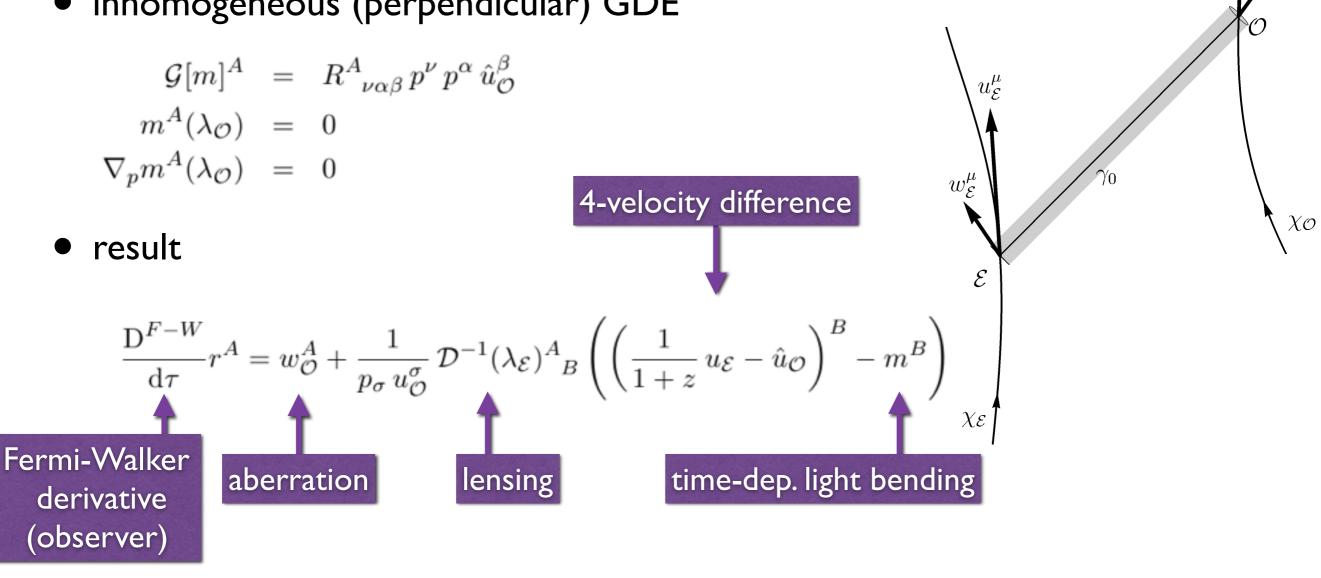
CosmoTorun [7

M. Korzyński, "Drift effects in cosmology"

 $u^{\mu}_{\mathcal{O}}$

Position drift

- parallel propagation of \hat{u}^{μ}_{O}
- inhomogeneous (perpendicular) GDE

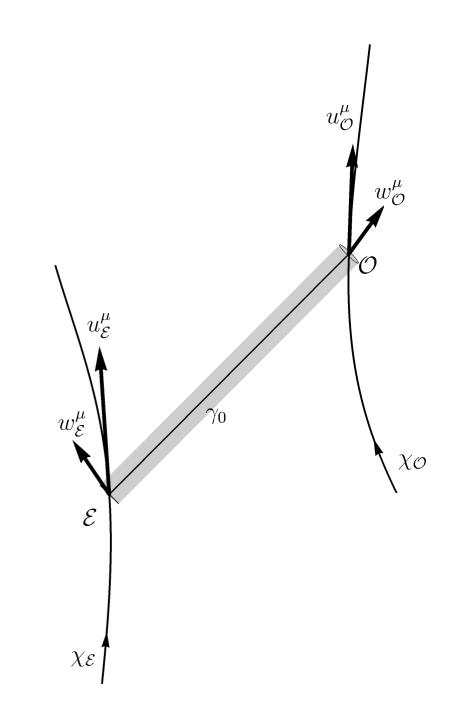


CosmoTorun 7

 $u^{\mu}_{\mathcal{O}}$



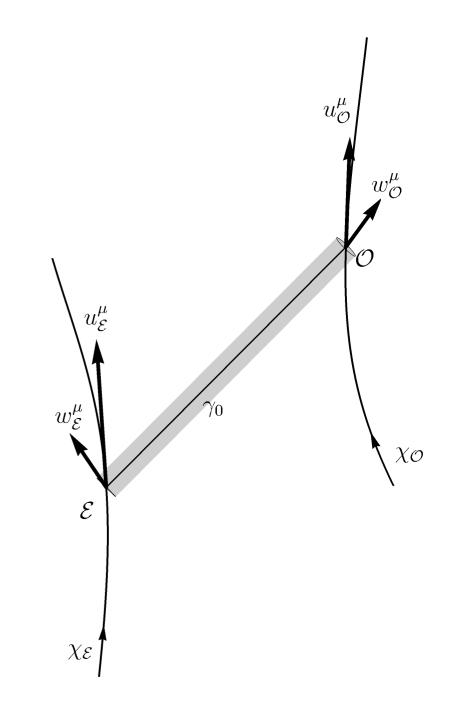
Redshift drift



CosmoTorun I7

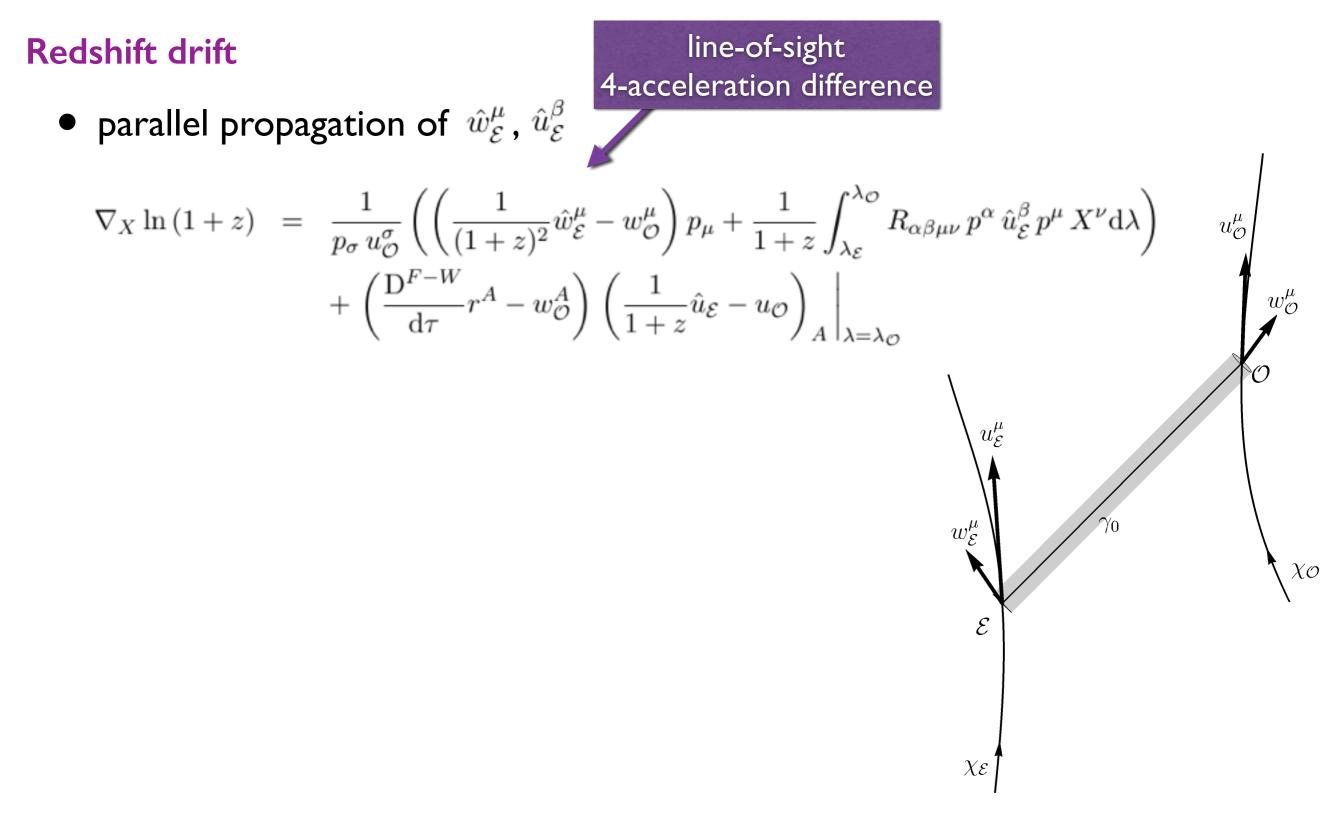
Redshift drift

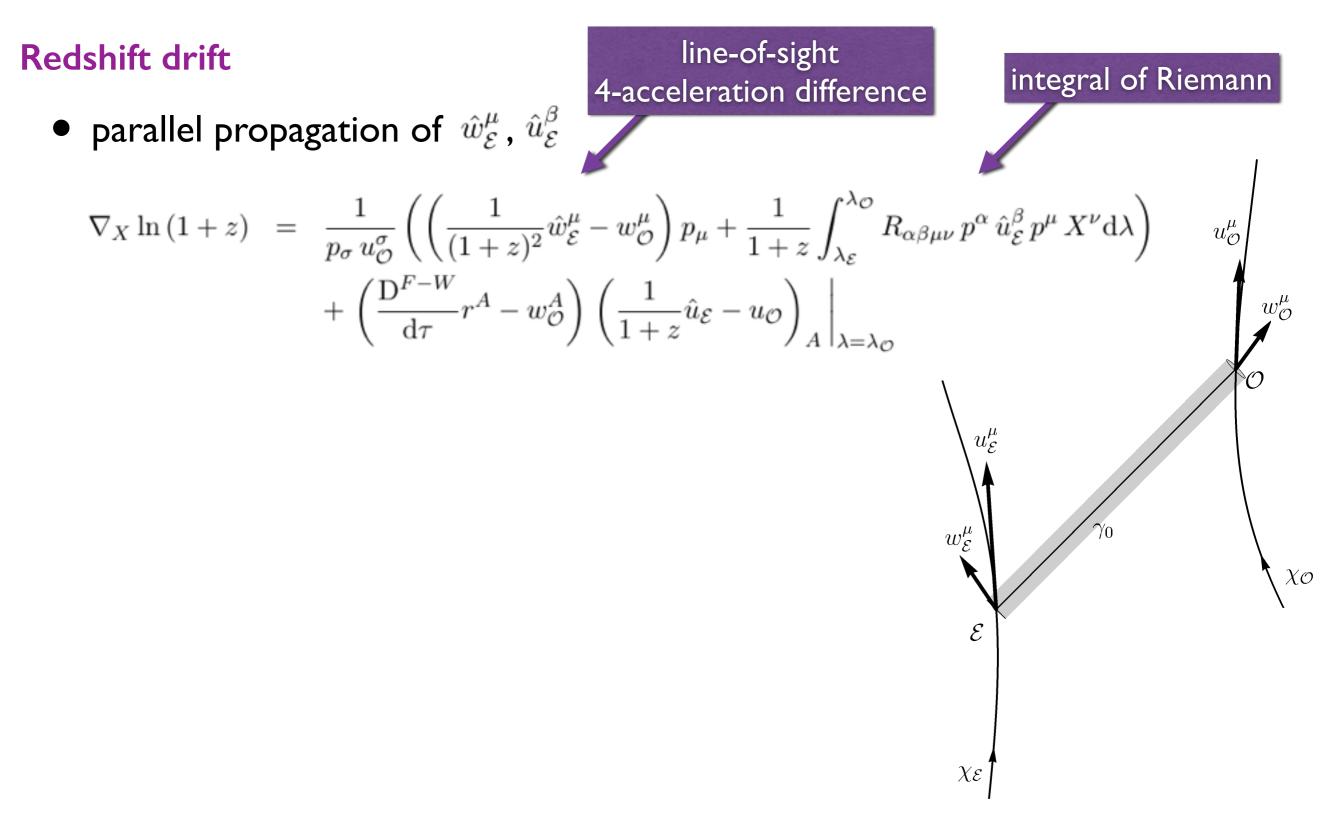
• parallel propagation of $\hat{w}^{\mu}_{\mathcal{E}}$, $\hat{u}^{\beta}_{\mathcal{E}}$

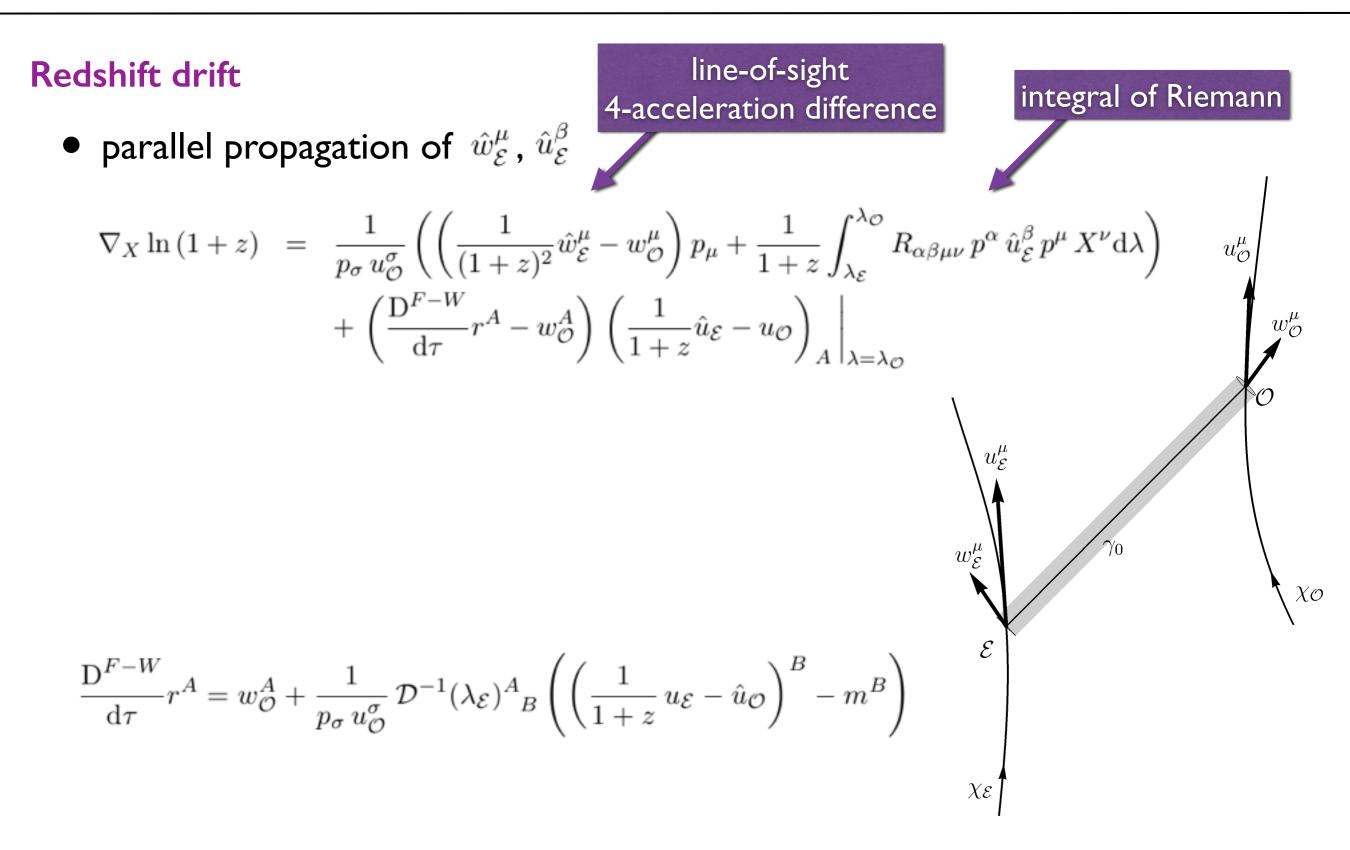


Redshift drift

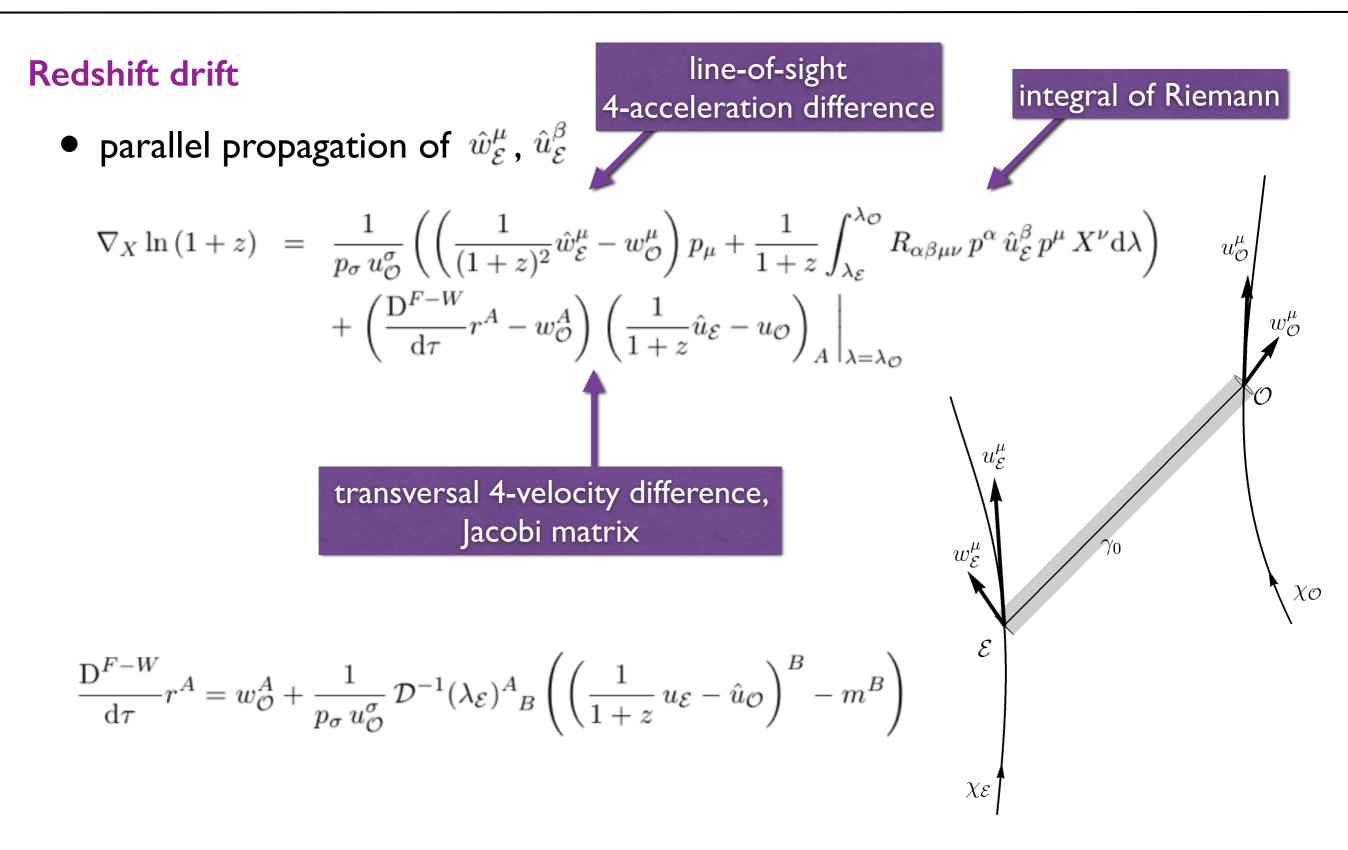
• parallel propagation of $\hat{w}^{\mu}_{\mathcal{E}}$, $\hat{u}^{\beta}_{\mathcal{E}}$





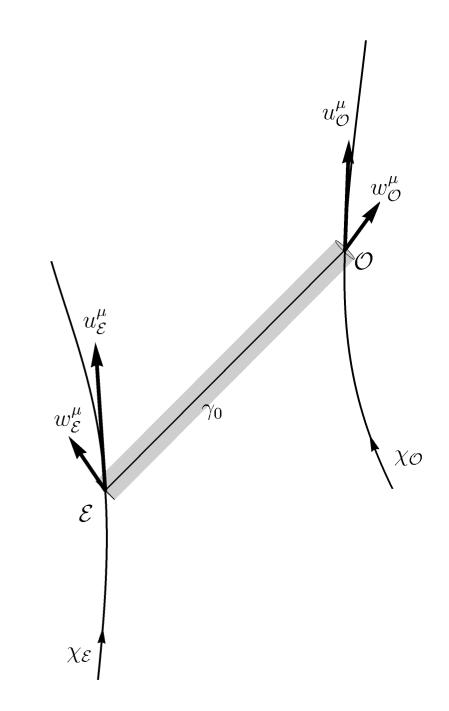


CosmoTorun I7





Jacobi matrix drift

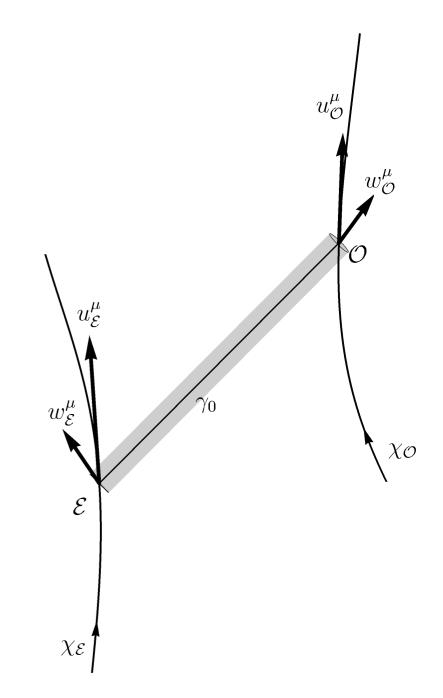


CosmoTorun I7

Jacobi matrix drift

• need to differentiate the GDE

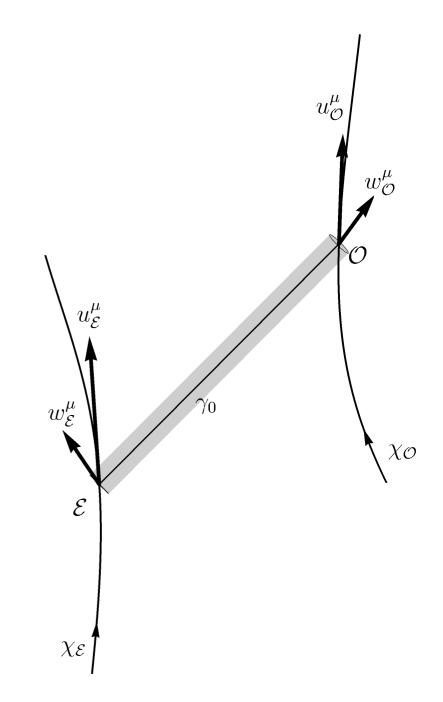
$$\mathcal{G}[\xi]^{\mu} \equiv \nabla_p \nabla_p \xi^{\mu} - R^{\mu}_{\ \nu\alpha\beta} \, p^{\nu} \, p^{\alpha} \, \xi^{\beta} = 0$$



Jacobi matrix drift

• need to differentiate the GDE

$$\nabla_X \qquad \mathcal{G}[\xi]^\mu \equiv \nabla_p \nabla_p \xi^\mu - R^\mu_{\ \nu\alpha\beta} \, p^\nu \, p^\alpha \, \xi^\beta = 0$$

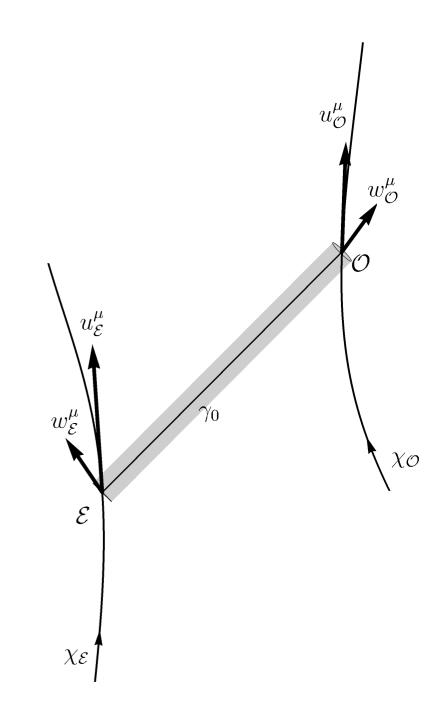


Jacobi matrix drift

• need to differentiate the GDE

$$\nabla_X \qquad \mathcal{G}[\xi]^{\mu} \equiv \nabla_p \nabla_p \xi^{\mu} - R^{\mu}_{\ \nu\alpha\beta} \, p^{\nu} \, p^{\alpha} \, \xi^{\beta} = 0$$

• inhomogeneous GDE



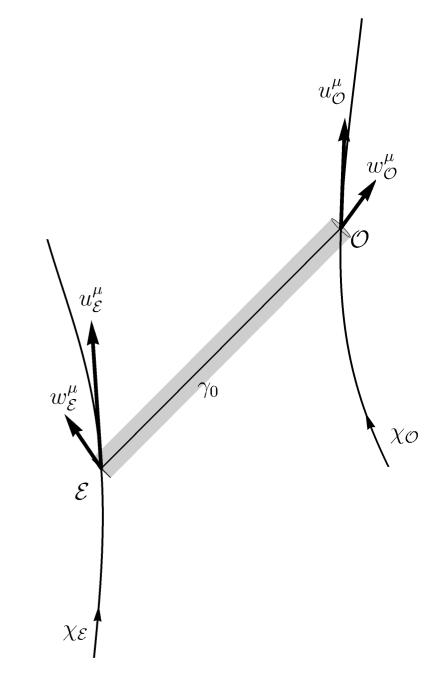
Jacobi matrix drift

• need to differentiate the GDE

$$\nabla_X \qquad \mathcal{G}[\xi]^{\mu} \equiv \nabla_p \nabla_p \xi^{\mu} - R^{\mu}_{\ \nu\alpha\beta} \, p^{\nu} \, p^{\alpha} \, \xi^{\beta} = 0$$

• inhomogeneous GDE

$$\nabla_p \nabla_p \left(\nabla_X \xi^\mu \right) - R^\mu_{\ \alpha\beta\nu} \, p^\alpha \, p^\beta \, \nabla_X \xi^\nu \quad = \quad \mathcal{M}^\mu_{\ \nu} \, \xi^\nu + \mathcal{N}^\mu_{\ \nu} \, \nabla_p \xi^\nu$$



Jacobi matrix drift

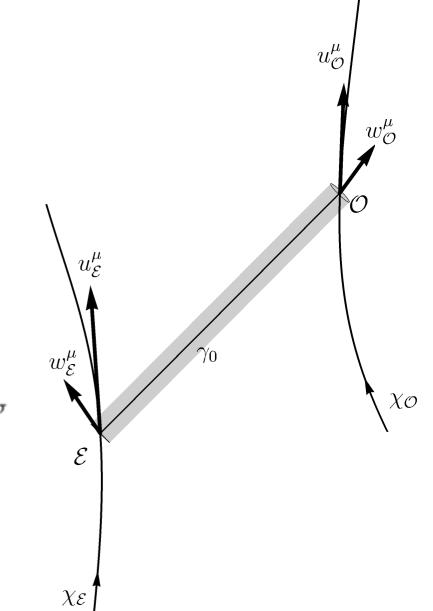
• need to differentiate the GDE

$$\nabla_X \qquad \mathcal{G}[\xi]^{\mu} \equiv \nabla_p \nabla_p \xi^{\mu} - R^{\mu}_{\ \nu\alpha\beta} \, p^{\nu} \, p^{\alpha} \, \xi^{\beta} = 0$$

• inhomogeneous GDE

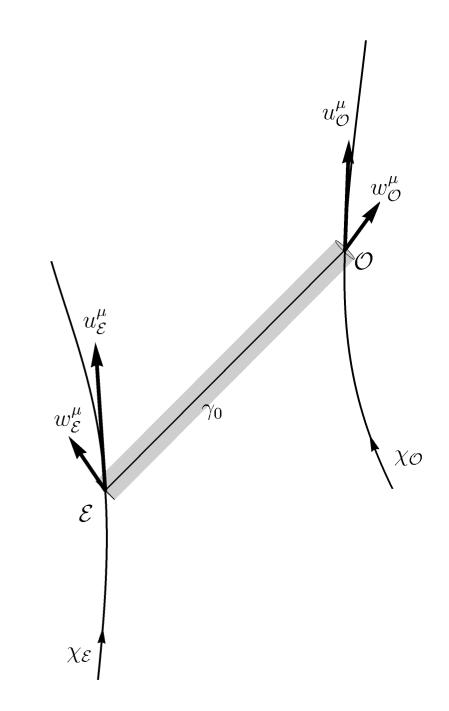
$$\nabla_p \nabla_p \left(\nabla_X \xi^\mu \right) - R^\mu_{\ \alpha\beta\nu} \, p^\alpha \, p^\beta \, \nabla_X \xi^\nu \quad = \quad \mathcal{M}^\mu_{\ \nu} \, \xi^\nu + \mathcal{N}^\mu_{\ \nu} \, \nabla_p \xi^\nu$$

$$\mathcal{M}^{\mu}_{\nu} = -\nabla_{\alpha} R^{\mu}_{\ \nu\rho\sigma} p^{\alpha} X^{\rho} p^{\sigma} + \nabla_{\alpha} R^{\mu}_{\ \beta\rho\nu} X^{\alpha} p^{\beta} p^{\rho} + 2R^{\mu}_{\ \beta\nu\sigma} \nabla_{p} X^{\beta} p^{\sigma}$$
$$\mathcal{N}^{\mu}_{\nu} = -2R^{\mu}_{\ \nu\rho\sigma} X^{\rho} p^{\sigma}$$





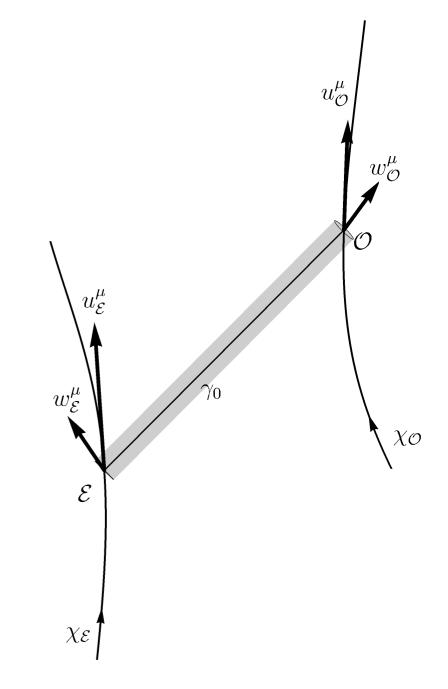
Jacobi matrix drift



CosmoTorun I7

Jacobi matrix drift

• vectors
$$f_1^{\ \mu}, f_2^{\ \mu}$$
 giving \mathcal{D}^A_B
 $\nabla_p \nabla_p f_A^{\ \mu} - R^{\mu}_{\ \alpha\beta\nu} p^{\alpha} p^{\beta} f_A^{\ \nu} = 0$
 $f_A^{\ \mu} (\lambda_{\mathcal{O}}) = 0$
 $\nabla_p f_A^{\ \mu} (\lambda_{\mathcal{O}}) = e_A^{\ \mu}$



Jacobi matrix drift

• vectors $f_1^{\ \mu}, f_2^{\ \mu}$ giving \mathcal{D}^{A_B}

$$\nabla_p \nabla_p f_A^{\ \mu} - R^{\mu}_{\ \alpha\beta\nu} p^{\alpha} p^{\beta} f_A^{\ \nu} = 0$$

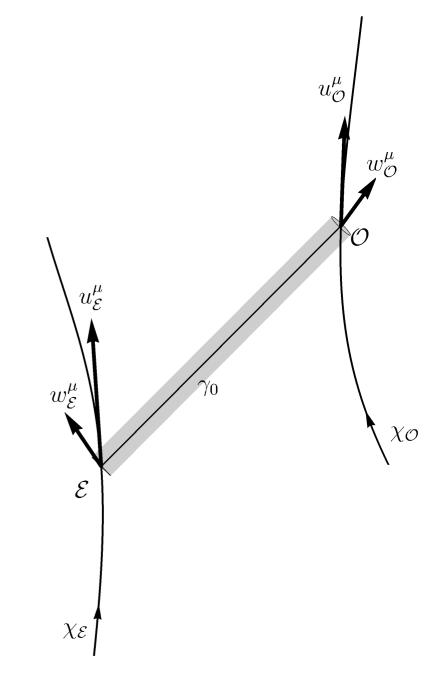
$$f_A^{\ \mu} (\lambda_{\mathcal{O}}) = 0$$

$$\nabla_p f_A^{\ \mu} (\lambda_{\mathcal{O}}) = e_A^{\ \mu}$$

• derivatives wrt X^{μ}

$$\nabla_p \nabla_p \left(\nabla_X f_A^{\ \mu} \right) - R^{\mu}_{\ \alpha\beta\nu} p^{\alpha} p^{\beta} \nabla_X f_A^{\ \mu} = \mathcal{M}^{\mu}_{\ \nu} f_A^{\ \nu} + \mathcal{N}^{\mu}_{\ \nu} \nabla_p f_A^{\ \nu}$$

+ initial data at O



Jacobi matrix drift

• vectors $f_1^{\ \mu}, f_2^{\ \mu}$ giving \mathcal{D}^{A_B}

$$\nabla_{p} \nabla_{p} f_{A}^{\ \mu} - R^{\mu}_{\ \alpha\beta\nu} p^{\alpha} p^{\beta} f_{A}^{\ \nu} = 0$$

$$f_{A}^{\ \mu} (\lambda_{\mathcal{O}}) = 0$$

$$\nabla_{p} f_{A}^{\ \mu} (\lambda_{\mathcal{O}}) = e_{A}^{\ \mu}$$

• derivatives wrt X^{μ}

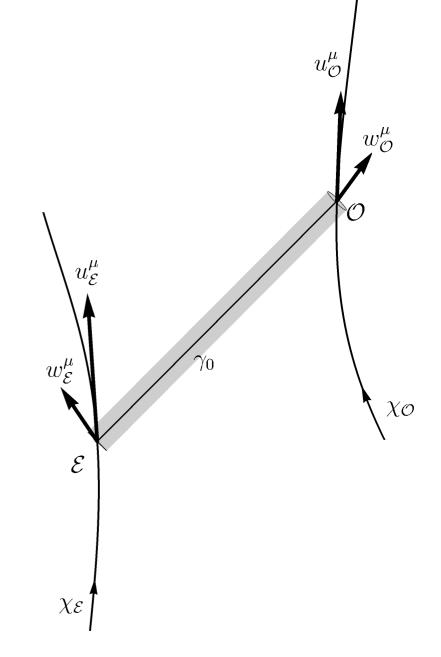
$$\nabla_p \nabla_p \left(\nabla_X f_A^{\ \mu} \right) - R^{\mu}_{\ \alpha\beta\nu} p^{\alpha} p^{\beta} \nabla_X f_A^{\ \mu} = \mathcal{M}^{\mu}_{\ \nu} f_A^{\ \nu} + \mathcal{N}^{\mu}_{\ \nu} \nabla_p f_A^{\ \nu}$$

+ initial data at O

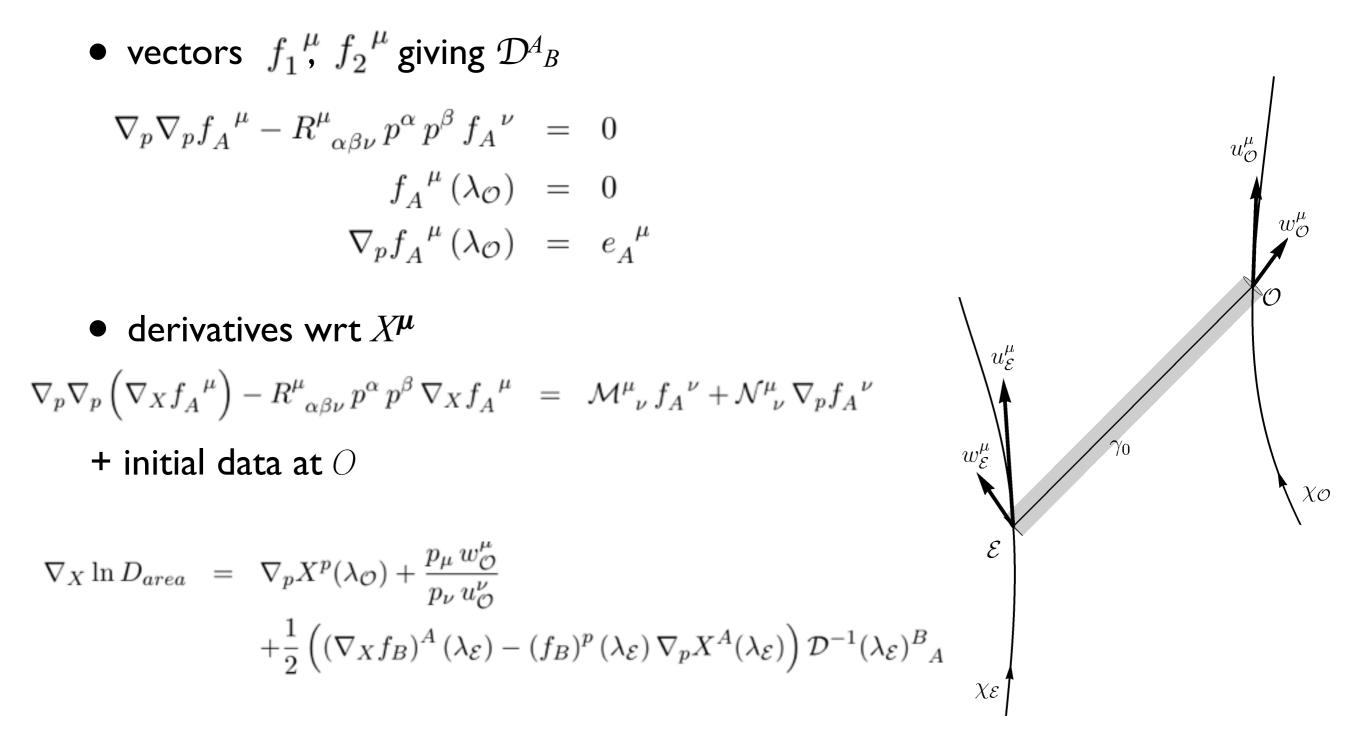
$$\nabla_X \mathcal{D}^A{}_B = (\nabla_X f_B)^A + \psi^A{}_C \ (f_B)^C - \nabla_p X^A \ (f_B)^p$$

antisymmetric,
irrelevant

CosmoTorun I7



Jacobi matrix drift



CosmoTorun I7

Jacobi matrix drift

• vectors $f_1^{\ \mu}, f_2^{\ \mu}$ giving \mathcal{D}^{A_B}

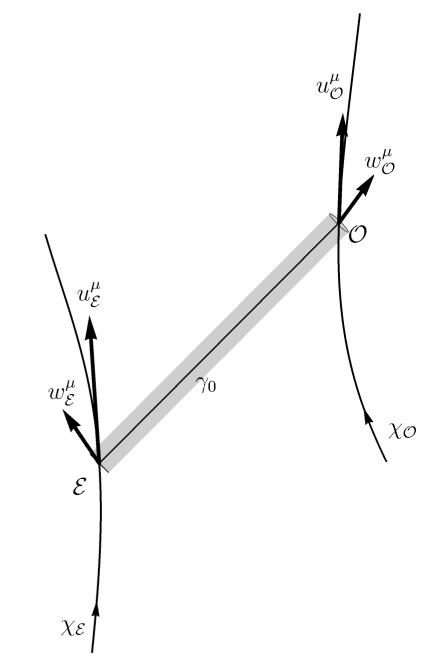
$$\nabla_{p} \nabla_{p} f_{A}^{\ \mu} - R^{\mu}_{\ \alpha\beta\nu} p^{\alpha} p^{\beta} f_{A}^{\ \nu} = 0$$
$$f_{A}^{\ \mu} (\lambda_{\mathcal{O}}) = 0$$
$$\nabla_{p} f_{A}^{\ \mu} (\lambda_{\mathcal{O}}) = e_{A}^{\ \mu}$$

• derivatives wrt X^{μ}

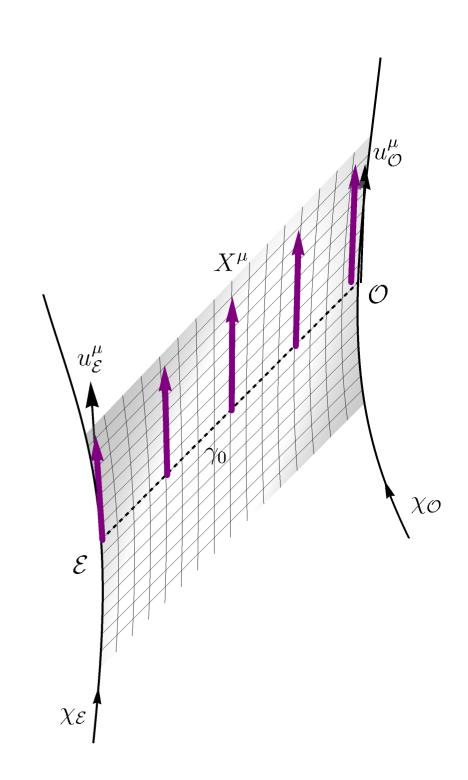
$$\nabla_p \nabla_p \left(\nabla_X f_A^{\ \mu} \right) - R^{\mu}_{\ \alpha\beta\nu} \, p^{\alpha} \, p^{\beta} \, \nabla_X f_A^{\ \mu} = \mathcal{M}^{\mu}_{\ \nu} \, f_A^{\ \nu} + \mathcal{N}^{\mu}_{\ \nu} \, \nabla_p f_A^{\ \nu}$$

+ initial data at ()

$$\nabla_X \ln D_{lum} = \nabla_X \ln D_{area} + 2\nabla_X \ln(1+z)$$



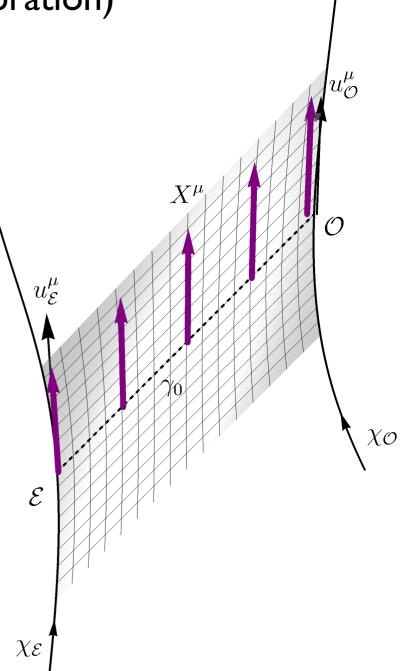
Remarks



CosmoTorun I7

Remarks

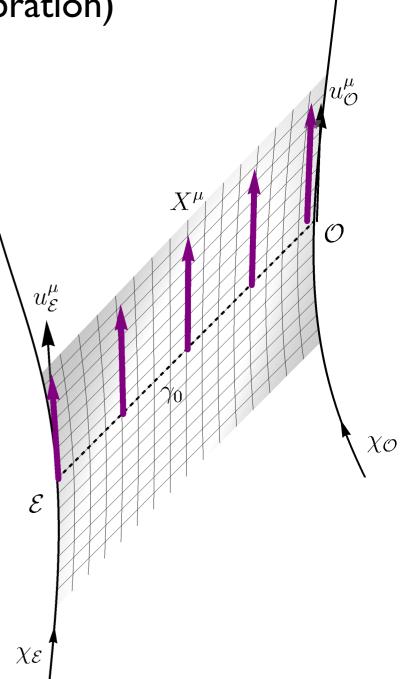
• $\nabla_X p^{\mu}$ at \mathcal{E} gives the viewing angle drift (apparent libration)



CosmoTorun17

Remarks

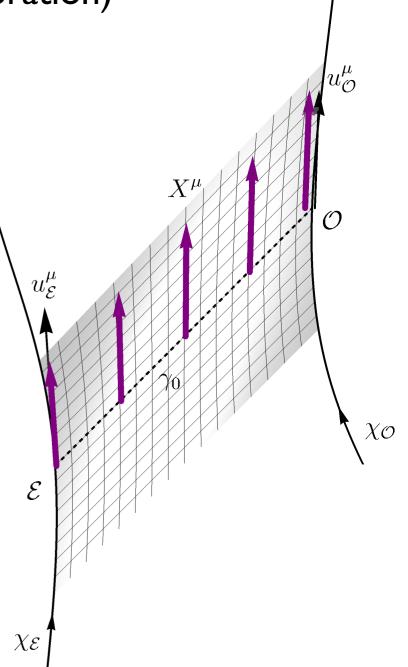
- $\nabla_X p^{\mu}$ at \mathcal{E} gives the viewing angle drift (apparent libration)
- caustics problematic



CosmoTorun I7

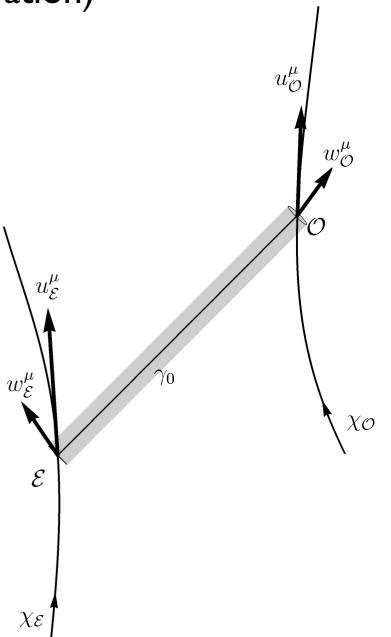
Remarks

- $\nabla_X p^{\mu}$ at \mathcal{E} gives the viewing angle drift (apparent libration)
- caustics problematic
- dependence on observer's and emitter's motion:



Remarks

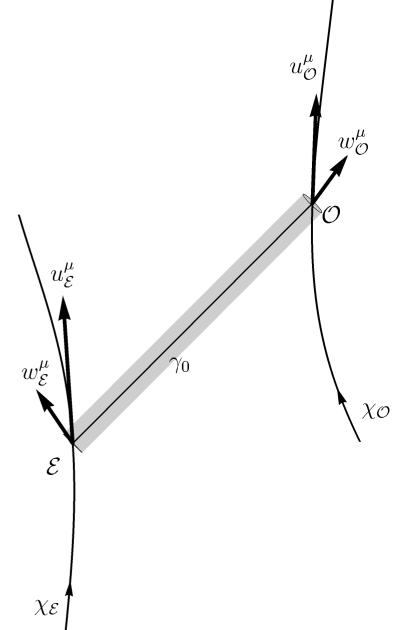
- $\nabla_X p^{\mu}$ at \mathcal{E} gives the viewing angle drift (apparent libration)
- caustics problematic
- dependence on observer's and emitter's motion:



Remarks

- $\nabla_X p^{\mu}$ at \mathcal{E} gives the viewing angle drift (apparent libration)
- caustics problematic
- dependence on observer's and emitter's motion:

$$\begin{split} \theta^{A} &\equiv \theta^{A} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu} \right) \\ z &\equiv z \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu} \right) \\ \mathcal{D}^{A}{}_{B} &\equiv \mathcal{D}^{A}{}_{B} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ D_{area} &\equiv D_{area} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ D_{lum} &\equiv D_{lum} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \theta^{A} &\equiv \dot{\theta}^{A} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \ln(1+z) &\equiv (1+z)^{-1} \dot{z} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \mathcal{D}^{A}{}_{B} &\equiv \dot{\mathcal{D}}^{A}{}_{B} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \ln D_{area} &\equiv (\ln D_{area}) \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \ln D_{lum} &\equiv (\ln D_{lum}) \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \end{split}$$

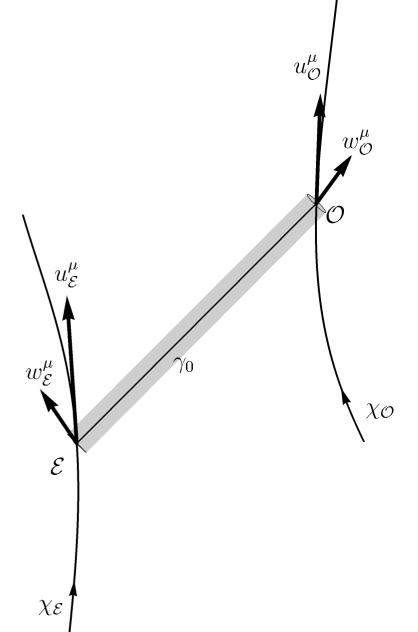


CosmoTorun I 7

Remarks

- $\nabla_X p^{\mu}$ at \mathcal{E} gives the viewing angle drift (apparent libration)
- caustics problematic
- dependence on observer's and emitter's motion:

$$\begin{split} \theta^{A} &\equiv \theta^{A} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu} \right) \\ z &\equiv z \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu} \right) \\ \mathcal{D}^{A}{}_{B} &\equiv \mathcal{D}^{A}{}_{B} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ D_{area} &\equiv D_{area} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ D_{lum} &\equiv D_{lum} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \theta^{A} &\equiv \dot{\theta}^{A} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \ln (1+z) &\equiv (1+z)^{-1} \dot{z} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \mathcal{D}^{A}{}_{B} &\equiv \dot{\mathcal{D}}^{A}{}_{B} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \ln D_{area} &\equiv (\ln D_{area})^{\cdot} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \ln D_{lum} &\equiv (\ln D_{lum})^{\cdot} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \end{split}$$

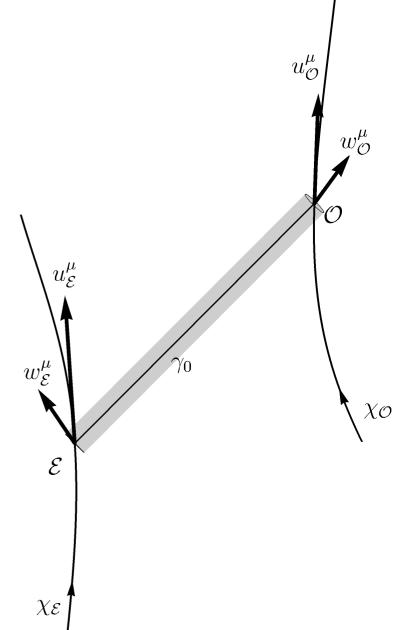


CosmoTorun I 7

Remarks

- $\nabla_X p^{\mu}$ at \mathcal{E} gives the viewing angle drift (apparent libration)
- caustics problematic
- dependence on observer's and emitter's motion:

$$\begin{split} \theta^{A} &\equiv \theta^{A} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu} \right) \\ z &\equiv z \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu} \right) \\ \mathcal{D}^{A}{}_{B} &\equiv \mathcal{D}^{A}{}_{B} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, y_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ D_{area} &\equiv D_{area} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, y_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ D_{lum} &\equiv D_{lum} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \theta^{A} &\equiv \dot{\theta}^{A} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \ln(1+z) &\equiv (1+z)^{-1} \dot{z} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \mathcal{D}^{A}{}_{B} &\equiv \dot{\mathcal{D}}^{A}{}_{B} \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \ln D_{area} &\equiv (\ln D_{area}) \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \\ \frac{d}{d\tau} \ln D_{lum} &\equiv (\ln D_{lum}) \left(\mathcal{O}, \mathcal{E}, u_{\mathcal{O}}^{\mu}, u_{\mathcal{E}}^{\mu}, w_{\mathcal{O}}^{\mu}, w_{\mathcal{E}}^{\mu}, R^{\mu}{}_{\nu\alpha\beta}, \nabla_{\gamma} R^{\mu}{}_{\nu\alpha\beta} \right) \end{split}$$



CosmoTorun I 7

Extension of the Sachs formalism

• Numerics: extension of ray tracing in numerical spacetimes. Given a null geodesic, observer and emitter obtain z, D_{ang} , D_{lum} and \mathcal{D}^{A}_{B} as well as \dot{z} , \dot{D}_{ang} , \dot{D}_{lum} , $\dot{\vartheta}^{A}$ and $\dot{\mathcal{D}}^{A}_{B}$ by solving ODE's

- Numerics: extension of ray tracing in numerical spacetimes. Given a null geodesic, observer and emitter obtain z, D_{ang} , D_{lum} and \mathcal{D}^{A}_{B} as well as \dot{z} , \dot{D}_{ang} , \dot{D}_{lum} , $\dot{\vartheta}^{A}$ and $\dot{\mathcal{D}}^{A}_{B}$ by solving ODE's
- Including the inhomogeneities into the theory of drift effects

- Numerics: extension of ray tracing in numerical spacetimes. Given a null geodesic, observer and emitter obtain z, D_{ang} , D_{lum} and \mathcal{D}^{A}_{B} as well as \dot{z} , \dot{D}_{ang} , \dot{D}_{lum} , $\dot{\vartheta}^{A}$ and $\dot{\mathcal{D}}^{A}_{B}$ by solving ODE's
- Including the inhomogeneities into the theory of drift effects
 - exact models

- Numerics: extension of ray tracing in numerical spacetimes. Given a null geodesic, observer and emitter obtain z, D_{ang} , D_{lum} and \mathcal{D}^{A}_{B} as well as \dot{z} , \dot{D}_{ang} , \dot{D}_{lum} , $\dot{\vartheta}^{A}$ and $\dot{\mathcal{D}}^{A}_{B}$ by solving ODE's
- Including the inhomogeneities into the theory of drift effects
 - exact models
 - standard perturbation theory

- Numerics: extension of ray tracing in numerical spacetimes. Given a null geodesic, observer and emitter obtain z, D_{ang} , D_{lum} and \mathcal{D}^{A}_{B} as well as \dot{z} , \dot{D}_{ang} , \dot{D}_{lum} , $\dot{\vartheta}^{A}$ and $\dot{\mathcal{D}}^{A}_{B}$ by solving ODE's
- Including the inhomogeneities into the theory of drift effects
 - exact models
 - standard perturbation theory
 - other approximations: local perturbation theory, PN formalism (Sanghai, Clifton 2015, 2016)...

- Numerics: extension of ray tracing in numerical spacetimes. Given a null geodesic, observer and emitter obtain z, D_{ang} , D_{lum} and \mathcal{D}^{A}_{B} as well as \dot{z} , \dot{D}_{ang} , \dot{D}_{lum} , $\dot{\vartheta}^{A}$ and $\dot{\mathcal{D}}^{A}_{B}$ by solving ODE's
- Including the inhomogeneities into the theory of drift effects
 - exact models
 - standard perturbation theory
 - other approximations: local perturbation theory, PN formalism (Sanghai, Clifton 2015, 2016)...
 - stochastic approach (Fleury et al 2013, Fleury 2014)

- Numerics: extension of ray tracing in numerical spacetimes. Given a null geodesic, observer and emitter obtain z, D_{ang} , D_{lum} and \mathcal{D}^{A}_{B} as well as \dot{z} , \dot{D}_{ang} , \dot{D}_{lum} , $\dot{\vartheta}^{A}$ and $\dot{\mathcal{D}}^{A}_{B}$ by solving ODE's
- Including the inhomogeneities into the theory of drift effects
 - exact models
 - standard perturbation theory
 - other approximations: local perturbation theory, PN formalism (Sanghai, Clifton 2015, 2016)...
 - stochastic approach (Fleury et al 2013, Fleury 2014)
- Tests of large scale homogeneity/isotropy

- Numerics: extension of ray tracing in numerical spacetimes. Given a null geodesic, observer and emitter obtain z, D_{ang} , D_{lum} and \mathcal{D}^{A}_{B} as well as \dot{z} , \dot{D}_{ang} , \dot{D}_{lum} , $\dot{\vartheta}^{A}$ and $\dot{\mathcal{D}}^{A}_{B}$ by solving ODE's
- Including the inhomogeneities into the theory of drift effects
 - exact models
 - standard perturbation theory
 - other approximations: local perturbation theory, PN formalism (Sanghai, Clifton 2015, 2016)...
 - stochastic approach (Fleury et al 2013, Fleury 2014)
- Tests of large scale homogeneity/isotropy
- Characterizations of the FLRW spacetime (Krasiński 2011)

Extension of the Sachs formalism

- Numerics: extension of ray tracing in numerical spacetimes. Given a null geodesic, observer and emitter obtain z, D_{ang} , D_{lum} and \mathcal{D}^{A}_{B} as well as \dot{z} , \dot{D}_{ang} , \dot{D}_{lum} , $\dot{\vartheta}^{A}$ and $\dot{\mathcal{D}}^{A}_{B}$ by solving ODE's
- Including the inhomogeneities into the theory of drift effects
 - exact models
 - standard perturbation theory
 - other approximations: local perturbation theory, PN formalism (Sanghai, Clifton 2015, 2016)...
 - stochastic approach (Fleury et al 2013, Fleury 2014)
- Tests of large scale homogeneity/isotropy
- Characterizations of the FLRW spacetime (Krasiński 2011)

to be published soon